Comparison of induced and spontaneous emission

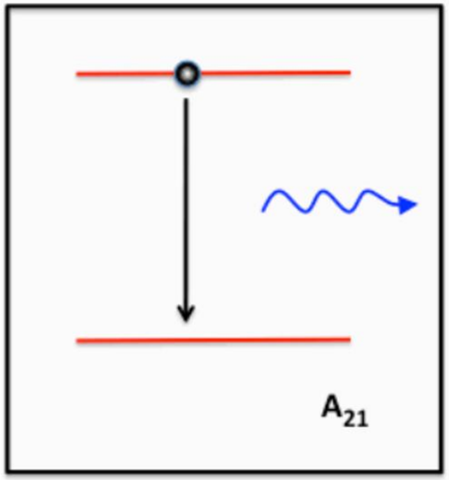
There was a home work:

- When (at what temperatures, wavelengths) is spontaneous or induced emission stronger?

  Assume LTE (blackbody)

Spontaneous & Stimulated emission

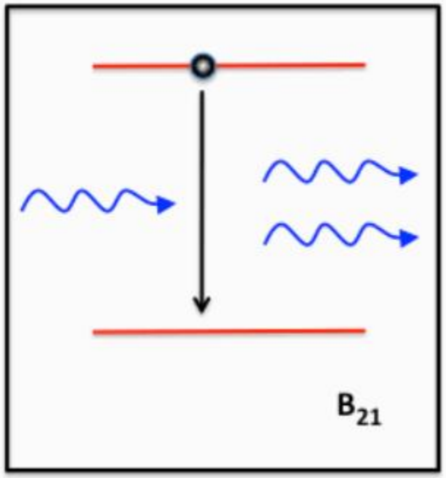
Spontaneous emission



Spontaneous emission

- The system goes from an upper level *u* to a lower level *l* spontaneously.
- Occurs independently of the radiation field.
- Emits **isotropically**.

Stimulated emission



Stimulated emission

- The system goes from an upper level *u* to a lower level *l* stimulated by the presence of a radiation field ($h\nu$ corresponding to the energy difference between levels *u* and *l*).
- Stimulated emission occurs into the **same** state (frequency, direction, polarization) as the photon that stimulated the emission.

Relation between Einstein coefficients

$$\frac{A_{21}}{B_{21}} = \frac{2h\nu_{21}^3}{c^2} \quad \rightarrow \quad A_{21} = B_{21}\frac{2h\nu_{21}^3}{c^2}$$

$$\frac{g_1 B_{12}}{g_2 B_{21}} = 1 \quad \rightarrow \quad g_1 B_{12} = g_2 B_{21}$$

Einstein's coefficients concern the probability that a particle spontaneously emits a photon, the probability to absorb a photon, and the probability to emit a photon under the influence of another incoming photon.
Einstein's coefficients are valid for all radiation fields.

Induced and Spontaneous emission

When is spontaneous emission stronger?

Total amount of emitted photons per unit time at a given frequency is

Spontaneous emission: $\eta_{sp}=n_2A_{21}$

Stimulated emission: $\eta_{st}=n_2B_{21}J_\nu$

$$A_{21} = B_{21}\frac{2h\nu_{21}^3}{c^2}$$

$$g_1B_{12} = g_2B_{21}$$

$$\frac{\eta_{sp}}{\eta_{st}} = \frac{n_2A_{21}}{n_2B_{21}J} = \frac{2h\nu_{21}^3}{c^2J}$$

TE: blackbody, J=$B_\nu(T)$

$$B_\nu(T) = \frac{2h\nu_{21}^3}{c^2}\left(e^{\frac{h\nu_{21}}{kT}} - 1\right)^{-1}$$

$$\frac{\eta_{sp}}{\eta_{st}} = e^{\frac{h\nu_{21}}{kT}} - 1$$

$$e^{\frac{h\nu_{21}}{kT}} \geq 2 \quad \Rightarrow \quad h\nu_{21} \geq kT\ln 2 \quad \Rightarrow \quad \lambda_* \leq \frac{hc}{kT\ln 2} = \frac{2.076\times 10^8}{T}\text{Å}$$

At wavelengths shorter than λ_* spontaneous emission is dominant

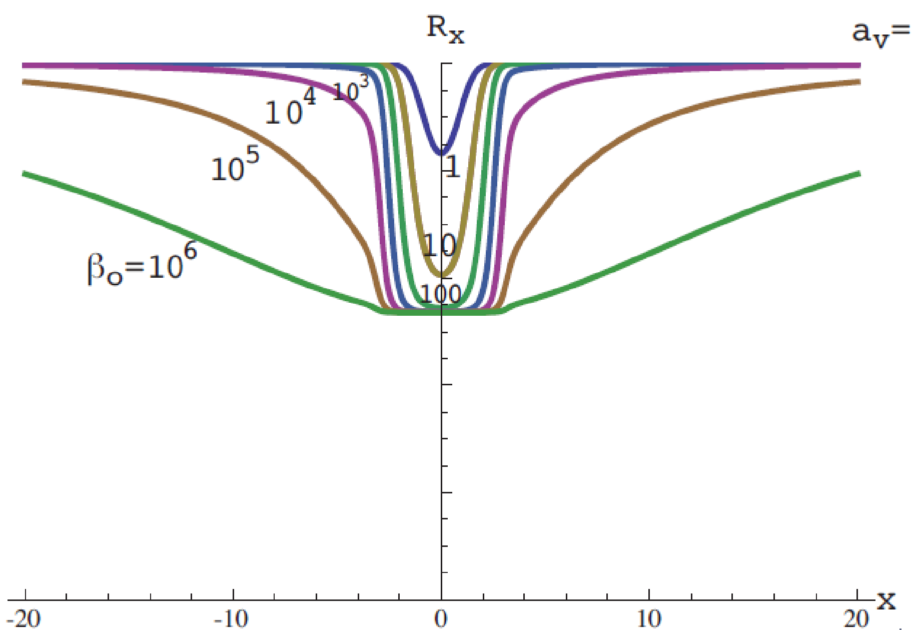
T=5777K → λ_*≈41000 Å λ_* =6563 Å → T≈31600K λ_* =4340Å → T≈48000K

Non-LTE

NON-LTE
STATISTICAL EQUILIBRIUM
TWO-LEVEL APPROXIMATION
THE LINE SOURCE FUNCTION
LTE VERSUS NON-LTE

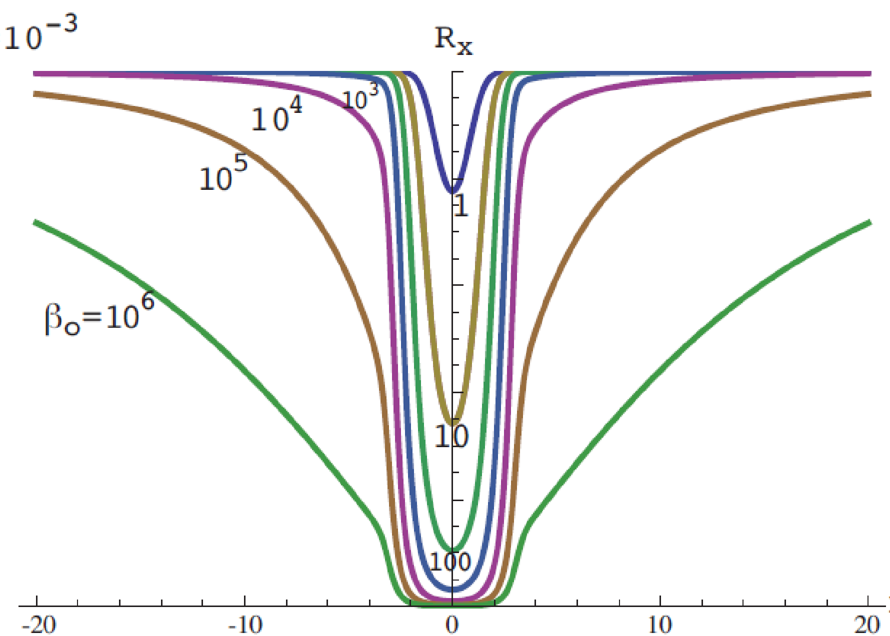
From the previous lecture

$\zeta = 1$ (LTE)



In LTE, the residual flux is non-zero even for strong absorption lines because we see the star surface with non-zero temperature.

$\zeta = 0$ (non-LTE)



In non-LTE, no photon emerges from surface due to scattering. Cores of strong scattering lines are **dark**!

LTE versus non-LTE?

- Most studies of stellar atmospheres are performed under LTE, where the thermodynamic state of the plasma is described via the Saha-Boltzmann equation as a function of local T and N_e. However, LTE strictly holds only deep in the interior when collisions dominate, and the photon mean-free-path is small.

- For a more accurate physical description, the non-local nature of the radiation field and its interaction with the plasma must be accounted for. This requires consideration of the detailed atomic processes for excitation and ionization, as expressed in the rate equations of statistical equilibrium (non-LTE case).

- Departure coefficients b = pop(non-LTE)/pop(LTE)

What does non-LTE mean?

The level populations of atoms are governed by the rates of all (collisional and radiative) processes, by which an atom leaves a certain state i to some other state j (if bound) or k (if unbound) and vice versa.

Bound-bound	Bound-free
RADIATIVE	
Photoabsorption (R_{ij})	Photoionization (R_{ik})
Spontaneous + stimulated emission (R_{ji})	Spontaneous + stimulated recombination (R_{ki})
COLLISIONAL	
Excitation (C_{ij})	Ionization (C_{ik})
De-excitation (C_{ji})	Recombination (C_{ki})

The total upward rate $P_{ij}=C_{ij}+R_{ij}$, whilst the total downward rate is $P_{ji}=C_{ji}+R_{ji}$

LTE vs NLTE

- LTE: population numbers follow Saha-Boltzmann Equation
  $n_i = n_i(T, n_e)$
- NLTE: population numbers depend on radiation field
  $n_i = n_i(T, n_e, J)$
- Need to take into account the sum of all processes that decrease and increase population for a given level *i*:

$$\frac{d}{dt} n_i = \sum_{j \neq i} n_j P_{ji} - n_i \sum_{j \neq i} P_{ij}$$

In stellar atmospheres typically:
$dn_i / dt = 0$ (stationary)

Complete rate equations

For each atomic level i of each ion, of each chemical element we have:

$$-n_i\left[\sum_{j>i}(R_{ij}+C_{ij})+\sum_{j<i}(R_{ij}+C_{ij})\right]+\sum_{j>i}n_j(R_{ji}+C_{ji})+\sum_{j<i}n_j(R_{ji}+C_{ji})=\frac{dn_i}{dt}$$

- $\sum_{j>i}(R_{ij}+C_{ij})$: excitation and ionization
- $\sum_{j<i}(R_{ij}+C_{ij})$: de-excitation and recombination

Together: rates out of i

- $\sum_{j>i}n_j(R_{ji}+C_{ji})$: de-excitation and recombination
- $\sum_{j<i}n_j(R_{ji}+C_{ji})$: excitation and ionization

Together: rates into i

In steady-state, $dn_i/dt=0$

Statistical equilibrium

- Statistical equilibrium, also known as rate equations:

$$\frac{dn_i}{dt} = \sum_{j\neq i}^{N} n_j P_{ji} - n_i \sum_{j\neq i}^{N} P_{ij} =$$

$$= \sum_{j\neq i}^{N} n_j \left(R_{ji} + C_{ji}\right) + n_p (R_{ki} + C_{ki}) - n_i \sum_{j\neq i}^{N} \left(R_{ij} + C_{ij}\right) - n_i (R_{ik} + C_{ik}) = 0$$

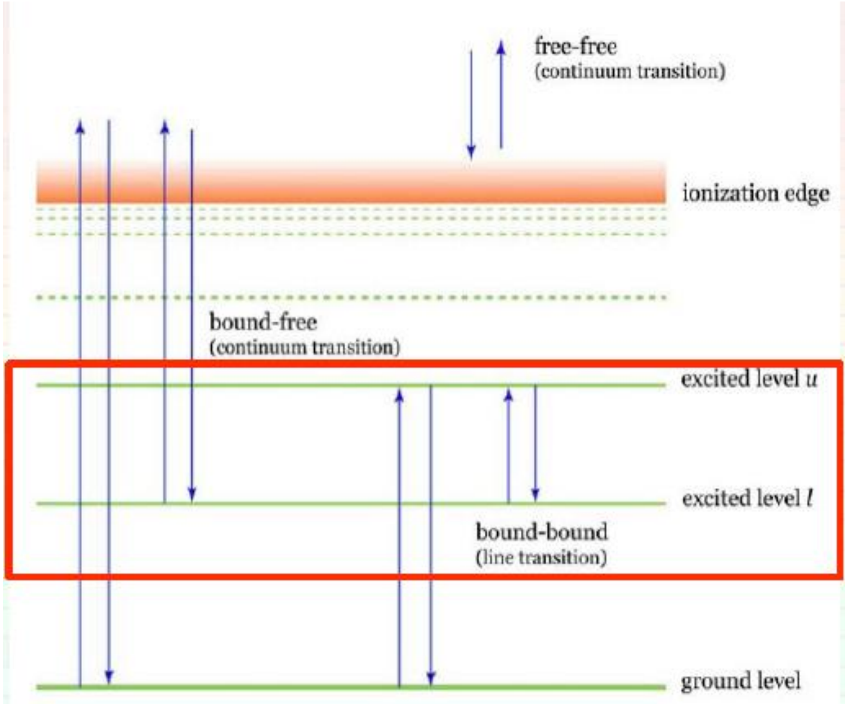
Lines Recombination Lines Ionization

- Particle conservation: $$\sum_{i=1}^{N} n_i = n_T$$

- By "non-LTE", we refer to the solution of these equations of statistical equilibrium or rate equations. This is **much** more challenging computationally than LTE…
- Rate equations represent a non-linear system of equations, we look for the solution vector via linearization, based on Newton-Raphson iteration.

Two-level approximation

Let's consider schematic line-formation cases with easy solution.



Consider an atomic model with only **two** important levels: lower l and upper u.

It is highly simplified:

not accurate, but provide insight into the mechanisms at work in real stellar atmospheres.
It well approximates the situation for some lines, *e.g.* resonance lines from the ground state.

Two-level approximation

Rate equations (statistical equilibrium: for a given level ***i*** the rate of transitions **out** = rate of transitions **in**)

$$n_i \sum_{j \neq i}^{N} (R_{ij} + C_{ij}) - n_i (R_{ik} + C_{ik}) = \sum_{j \neq i}^{N} n_j (R_{ji} + C_{ji}) + n_p (R_{ki} + C_{ki})$$

Lines — Ionization — Lines — Recombination

Consider two levels *u* and *l*: **isolating** the transitions between them:

$$n_l (R_{lu} + C_{lu}) + n_l \sum_{j \neq l,u} (R_{lj} + C_{lj}) + n_l (R_{lk} + C_{lk}) = n_u (R_{ul} + C_{ul}) + \sum_{j \neq l,u} n_j (R_{jl} + C_{jl}) + n_p (R_{kl} + C_{kl})$$

and **neglecting** all transitions involving $j \neq l,u$, plus recombinations/ionizations:

$$\Rightarrow \quad n_l (R_{lu} + C_{lu}) = n_u (R_{ul} + C_{ul})$$

Two-level approximation

$$n_l(R_{lu} + C_{lu}) = n_u(R_{ul} + C_{ul})$$

Einstein coefficients:
$n_l B_{lu} J_\nu = n_u A_{ul} + n_u B_{ul} J_\nu$

substituting for the *R* coefficients:

$$n_l\left(B_{lu}\int_0^\infty \varphi_\nu J_\nu d\nu + C_{lu}\right) = n_u\left(A_{ul} + B_{ul}\int_0^\infty \varphi_\nu J_\nu d\nu + C_{ul}\right)$$

Transition Probabilities: Radiative Processes

$$R_{ij} = B_{ij}\int_0^\infty \varphi_{ij}(\nu)\ J_\nu\ d\nu$$

Absorption

$$R_{ji} = A_{ji} + B_{ji}\int_0^\infty \varphi_{ij}(\nu)\ J_\nu\ d\nu$$

Spontaneous and stimulated Emission

Two-level approximation

$$n_l(R_{lu} + C_{lu}) = n_u(R_{ul} + C_{ul})$$

Einstein coefficients:
$n_l B_{lu} J_\nu = n_u A_{ul} + n_u B_{ul} J_\nu$

substituting for the *R* coefficients:

$$n_l\left(B_{lu}\int_0^\infty \varphi_\nu J_\nu d\nu + C_{lu}\right) = n_u\left(A_{ul} + B_{ul}\int_0^\infty \varphi_\nu J_\nu d\nu + C_{ul}\right)$$

Sanity check: LTE case

assuming collision rates dominate over radiative rates $\quad n_l C_{lu} = n_u C_{ul}$

remembering that $\quad C_{lu} = \left(\frac{n_u}{n_l}\right)^* C_{ul} = \frac{g_u}{g_l} e^{-E_{ul}/kT}\, C_{ul}$

$$\Rightarrow \quad n_l \frac{g_u}{g_l} e^{-E_{ul}/kT} C_{ul} = n_u C_{ul}$$

$$\Rightarrow \quad \frac{n_u}{n_l} = \frac{g_u}{g_l} e^{-E_{ul}/kT} = \left(\frac{n_u}{n_l}\right)^{\mathrm{LTE}}$$

Calculation of the line source function

$$\alpha_\nu^{line} = (n_l B_{lu} - n_u B_{ul}) J_\nu \qquad \varepsilon_\nu^{line} = n_u A_{ul} J_\nu$$

$$S_\nu^{line} = \frac{\varepsilon_\nu^{line}}{\alpha_\nu^{line}} = \frac{n_u A_{ul}}{n_l B_{lu} - n_u B_{ul}} = \frac{A_{ul}}{\frac{n_l}{n_u} B_{lu} - B_{ul}}$$

Einstein Coefficients:

$$B_{lu} = \frac{g_u}{g_l} B_{ul} \qquad A_{ul} = \frac{2h\nu^3}{c^2} B_{ul}$$

$$S_\nu^{\text{line}} = \frac{2h\nu^3}{c^2} \frac{1}{\frac{n_l g_u}{n_u g_l} - 1}$$

Note: this is the general expression for the line source function in **NLTE**.
We have not made use of any equilibrium condition.
It is always valid (not only in 2-level approximation).
What is different in the general case, is how n_l and n_u are computed.

Calculation of the line source function

If we substitute $$\frac{n_u}{n_l} = \frac{g_u}{g_l} e^{-E_{ul}/kT} = \left(\frac{n_u}{n_l}\right)^*$$ $$E = h\nu$$

we recover the Planck function $$S_\nu^{\rm line} = \frac{2h\nu^3}{c^2} \frac{1}{e^{h\nu/kT} - 1} = B_\nu(T)$$

For the 2-level atom we found $$n_l (B_{lu} \int_0^\infty \varphi_\nu J_\nu d\nu + C_{lu}) = n_u (A_{ul} + B_{ul} \int_0^\infty \varphi_\nu J_\nu d\nu + C_{ul})$$

$$\frac{n_l}{n_u} = \frac{1 + \frac{c^2}{2h\nu^3} \int \varphi_{\nu'} J_{\nu'} d\nu' + C_{ul}/A_{ul}}{\frac{g_u}{g_l} \left[\frac{c^2}{2h\nu^3} \int \varphi_{\nu'} J_{\nu'} d\nu' + e^{-h\nu/kT} C_{ul}/A_{ul}\right]}$$

Calculation of the line source function

Substituting n_l/n_u in S_ν:

$$S_\nu^{\text{line}} = \frac{2h\nu^3}{c^2}\frac{1}{\frac{n_l g_u}{n_u g_l} - 1}$$

$$S_\nu^{\text{line}} = \frac{2h\nu^3}{c^2}\frac{\frac{c^2}{2h\nu^3}\int \varphi_{\nu'} J_{\nu'} d\nu' + e^{-h\nu/kT} C_{ul}/A_{ul}}{1 + \frac{C_{ul}}{A_{ul}}(1 - e^{-h\nu/kT})} =$$

$$= \underbrace{\frac{1}{1 + \frac{C_{ul}}{A_{ul}}(1 - e^{-h\nu/kT})}}_{1-\varepsilon} \int \varphi_{\nu'} J_{\nu'} d\nu' + \underbrace{\frac{2h\nu^3}{c^2}\frac{e^{-h\nu/kT} C_{ul}/A_{ul}}{1 + \frac{C_{ul}}{A_{ul}}(1 - e^{-h\nu/kT})}}$$

$$\underbrace{\frac{2h\nu^3}{c^2}\frac{1}{e^{h\nu/kT} - 1}}_{B_\nu(T)} \underbrace{\frac{(1 - e^{-h\nu/kT}) C_{ul}/A_{ul}}{1 + \frac{C_{ul}}{A_{ul}}(1 - e^{-h\nu/kT})}}_{:= \varepsilon}$$

The line source function (1)

scattering term → thermal term

$$S_\nu^{\text{line}} = (1-\epsilon)\int_0^\infty \varphi_{\nu'} J_{\nu'} d\nu' + \epsilon\, B_\nu(T) = \frac{\int_0^\infty \varphi_{\nu'} J_{\nu'} d\nu' + \epsilon'\; B_\nu(T)}{1+\epsilon'}$$

$$\epsilon := \frac{(1-e^{-h\nu/kT})\; C_{ul}/A_{ul}}{1+\frac{C_{ul}}{A_{ul}}\;(1-e^{-h\nu/kT})} = \frac{\epsilon'}{1+\epsilon'}$$

destruction probability

Photons are either destroyed into thermal pool or scattered photons are created in thermal processes

From the previous lecture:

$$\boldsymbol{S(\tau) = \zeta B + (1-\zeta)\, J(\tau)}$$

absorption fraction $\zeta \equiv \dfrac{\alpha_{\text{abs}}}{\alpha_{\text{abs}} + \alpha_{\text{sc}}}$

Now we obtained that Line source function has similar terms except that we also allow for non-coherent scattering

The line source function (2)

$$\epsilon := \frac{\left(1 - e^{-h\nu/kT}\right) C_{ul}/A_{ul}}{1 + \frac{C_{ul}}{A_{ul}} \left(1 - e^{-h\nu/kT}\right)} = \frac{\epsilon'}{1+\epsilon'}$$

$$S_\nu^{\text{line}} = (1-\epsilon) \int_0^\infty \varphi_{\nu'} J_{\nu'} d\nu' + \epsilon \, B_\nu(T) = \frac{\int_0^\infty \varphi_{\nu'} J_{\nu'} d\nu' + \epsilon' \, B_\nu(T)}{1+\epsilon'}$$

Deep layers: collisions dominate → ε' >> 1 or ε = 1 thermal term dominant

$C_{ul} \gg A_{ul}$ $\varepsilon = 1$ → $S_\nu = B_\nu(T)$ **LTE**

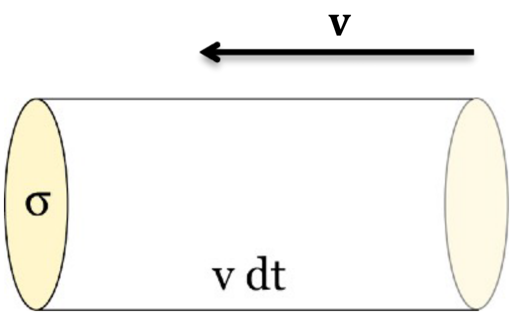
Higher layers: collisions non-important → ε'≈0 or ε = 0 scattering term dominant

$C_{ul} \ll A_{ul}$ $\varepsilon \approx 0$ → $S_\nu = \int \varphi_{\nu'} J_{\nu'} d\nu'$ **Extreme non-LTE**

From the previous lecture: $S_\nu = J_\nu$ for pure *coherent* scattering
now $S_\nu = \int \varphi_{\nu'} J_{\nu'} d\nu'$ *non-coherent* scattering

Transition probabilities: collisions

- probability of collision between atom/ion (cross section σ) and colliding particles in time dt: ~ σ v dt



collision cylinder: all particles in that volume collide with target atom

- rate of collisions = flux of colliding particles relative to atom/ion $(n_{coll}\, v)$ × cross section σ.

- for excitations: $$C_{ij} = n_{coll} \int_0^\infty \sigma_{ij}^{\mathrm{coll}}(v) v f(v) dv$$

- in a hot plasma free electrons dominate: $n_{coll} = n_e$

- f(v) dv is Maxwellian velocity distribution in stellar atmospheres established by e-e collisions.

Collisional rates (1)

For an electron with kinetic energy E exciting an atom

- Excitation:

$$C_{ij} = n_e \int_{E_{ij}}^{\infty} \sigma_{ij}(u) u f(u) du = n_e \int_{E_{ij}}^{\infty} \sigma_{ij}(E) \sqrt{\frac{2E}{m}} f(E) dE \propto \frac{n_e}{T^{3/2}} \int_{E_{ij}}^{\infty} \sigma_{ij}(E) e^{-E/kT} E\, dE \propto \frac{n_e}{T^{1/2}} e^{-E_{ij}/kT}$$

$$\sigma_{ij}(E) \propto 1/E$$

T is the kinetic temperature

- De-excitation:

$$C_{ji} = n_e \int_{0}^{\infty} \sigma_{ji}(u) u f(u) du = n_e \int_{0}^{\infty} \sigma_{ji}(E) \sqrt{\frac{2E}{m}} f(E) dE$$

where $f(E)$ is the (Maxwellian) energy distribution of the colliding particles.

- In TE, the principle of detailed balance gives

$$n_i C_{ij} = n_j C_{ji} \Rightarrow \frac{C_{ij}}{C_{ji}} = \frac{n_j}{n_i} = \frac{g_j}{g_i} e^{-E_{ij}/kT}$$

- Even if there is no TE, but we have a Maxwellian velocity distribution $\frac{C_{ij}}{C_{ji}} = \frac{g_j}{g_i} e^{-E_{ij}/kT}$

$$\frac{C_{ul}}{C_{lu}} = \frac{g_u}{g_l} e^{-(E_u - E_l)/kT}$$

Collisional rates (2)

$$\frac{C_{ul}}{C_{lu}} = \frac{g_u}{g_l} e^{-(E_u - E_l)/kT}$$

- Thus, if the excitations and de-excitations are due to collisions, the occupation numbers follow the Boltzmann formula for the kinetic temperature.
- We can conclude that in gases with high enough densities to make collisional excitations and de-excitations more important than the radiative processes, the occupation numbers follow the Boltzmann formula for the kinetic temperature.
- This means that the excitation temperature equals the kinetic temperature, which in turn means that the source function equals the Planck function for the kinetic temperature, which means **we have LTE**.

Two-level approximation

- Moving outward in the photosphere scattering term dominates.
- At some point we reach the region where photons are being lost from the star (small optical depth)

→ J_ν decreases with height
→ S_ν decreases with height
→ absorption line

- line absorption coefficient larger at line center → see higher layers
- wings form in deeper layers than line core

Wing can form in LTE conditions whilst a line core in non-LTE

- 2-level atom is a special NLTE case
- In general, the coupling between J_ν, n_i and S_ν is far more complicated

NLTE: Occupation numbers (1)

We obtain a system of linear equations for n_i:

$$A \cdot \begin{pmatrix} n_1 \\ n_2 \\ \cdots \\ n_p \end{pmatrix} = \mathbf{X}$$

Where matrix A contains terms:

$$\int_0^\infty \varphi_{ij}(\nu) \int_{4\pi} I_\nu(\omega) \frac{d\omega}{4\pi} d\nu$$

combine with equation of transfer:

$$\mu \frac{dI_\nu(\omega)}{dr} = -\kappa_\nu I_\nu(\omega) + \epsilon_\nu$$

$$\kappa_\nu = \sum_{i=1}^{N} \sum_{j=i+1}^{N} \sigma_{ij}^{\text{line}}(\nu) \left(n_i - \frac{g_i}{g_j} n_j\right) + \sum_{i=1}^{N} \sigma_{ik}(\nu) \left(n_i - n_i^* e^{-h\nu/kT}\right) + n_e n_p \sigma_{kk}(\nu, T) \left(1 - e^{-h\nu/kT}\right) + n_e \sigma_e$$

$$\epsilon_\nu = \ldots$$

non-linear system of integro-differential equations

NLTE: Occupation numbers (2)

Iteration required:

radiative processes depend on radiation field

$$R_{ij} = B_{ij} \int_0^\infty \varphi_{ij}(\nu)\ J_\nu\, d\nu$$

↓

radiation field depends on opacities

$$\mu \frac{dI_\nu(\omega)}{dr} = -\kappa_\nu I_\nu(\omega) + \epsilon_\nu$$

↓

opacities depend on occupation numbers

$$\kappa_\nu^{\mathrm{b-f}} = n_l\ \sigma_{lk}(\nu)$$

requires database of atomic quantities: energy levels, transitions, cross sections

20...1000 levels per ion – 3-5 ionization stages per species – » 30 species

➔ fast algorithm to calculate radiative transfer required

LTE

LTE is a **good** approximation, if:

1) Collisional rates **dominate** for all transitions

$R_{ij} << C_{ij}$ so $P_{ij}\ (= R_{ij} + C_{ij}) \sim C_{ij}$

Since $C_{ij}/C_{ji} = (n_i/n_j)^*$

Solution of rate equations -> LTE

2) $J_\nu = B_\nu$ is a good approximation at all frequencies

$n_i R_{ij} = n_j R_{ji}$ so $n_i/n_j = (n_i/n_j)^*$

Solution of rate equations -> LTE

Non-LTE

LTE is a bad approximation, if:

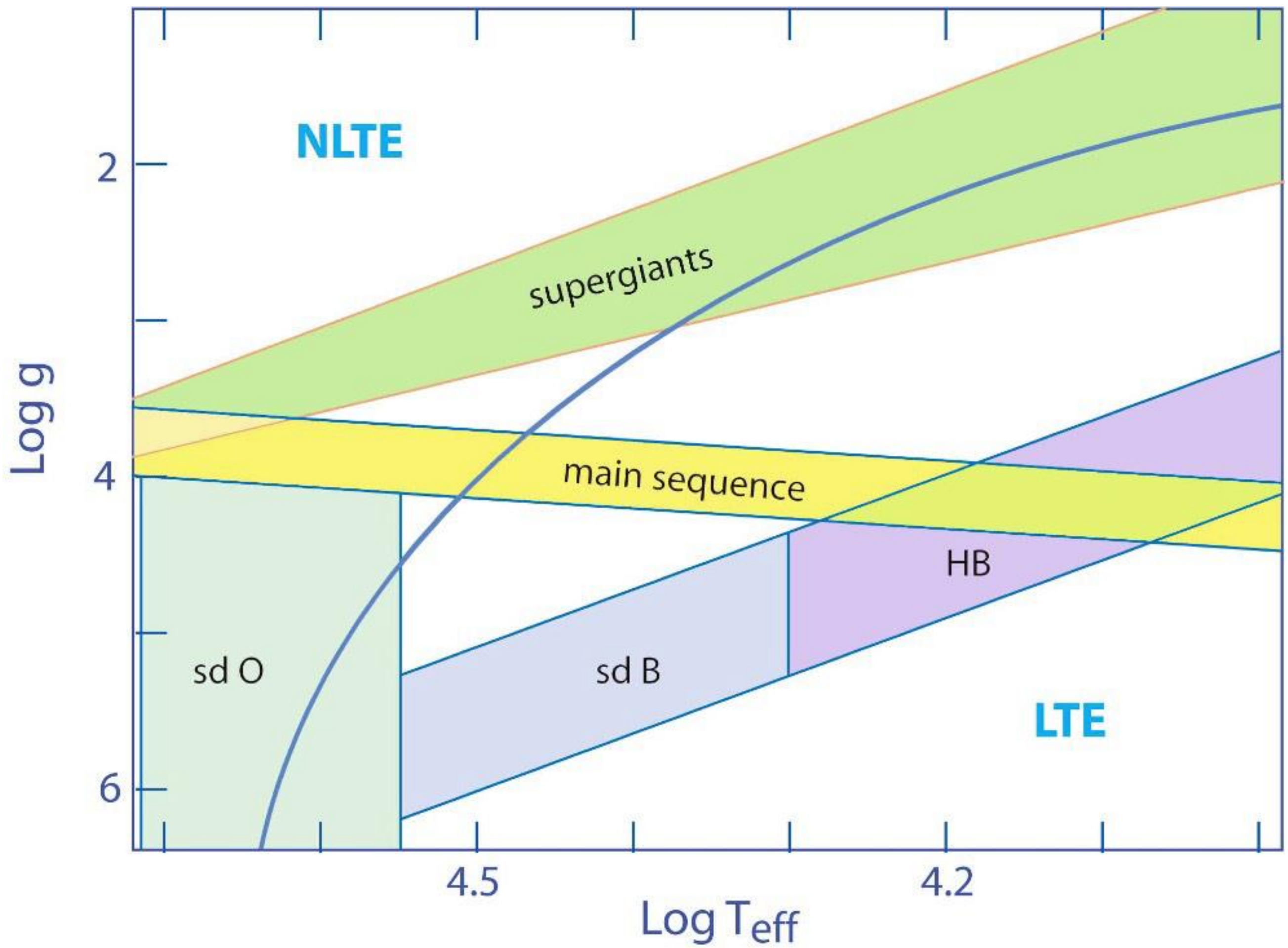
1) Collisional rates are small
2) Radiative rates are large
3) Mean free path of photons is larger than that of electrons

- Large deviations from LTE may be expected for low density gas in which the radiation field deviates strongly from the Planck function for the kinetic temperature.
- Non-LTE needs to be considered for
  (a) hot stars, whose atmospheres are rapidly expanding
  (b) low density chromospheres and coronae of Solar-type stars
  (c) low T_{eff} of very cool stars (in which electron densities are low)
  (d) nebulae
  (e) ISM

Non-LTE effects in scattering

- Deep in the atmosphere, collisions are frequent, radiation field is close to Planck and populations follow Boltzmann law. → LTE.
- Close to the boundary, radiation can escape freely, density drops, collisional rates decrease, radiative rates are not enough to populate upper levels.
  As a result, the upper level can be underpopulated.
  Therefore, the source function deviates from Planck function.
- Even if the only scattering (no true absorption) occurs in the atmosphere, an absorption line forms.

LTE vs NLTE

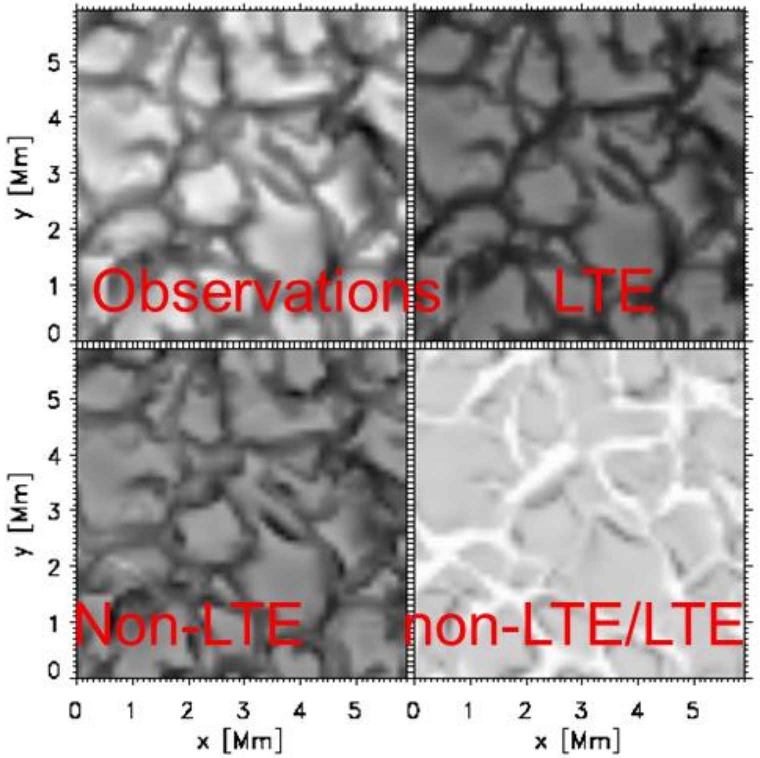

NLTE
supergiants
main sequence
HB
sd O
sd B
LTE
2
4
6
Log g
4.5
4.2
Log Teff

Non-LTE in the Sun

- Chromosphere & corona in non-LTE, since the radiation field corresponds to a diluted Planck function for the effective temperature of the Sun, whilst the kinetic temperature in the coronae may be several 10^6 K.
- Photospheric departures from LTE occur.
  Weak lines of low-abundance species often show departures from LTE (e.g. they reverse to emission lines on the solar disc just inside the limb). Cores of strong lines may depart from LTE, while the wings may remain in LTE.
- Non-LTE is most relevant in the Solar context via inaccuracies in elemental abundances obtained with the LTE assumption (typically 0.05 dex), although effect is greatest from comparison between latest 3D vs earlier 1D models.

Solar Oxygen abundance

- Until recently, commonly adopted Solar oxygen abundance was log(O/H)+12=8.93 suggested by analyses of [OI] 6300Å (Lambert 1978) and OH lines in IR using 1D LTE models.
- Asplund et al. (2004) has used 3D analyses of the [OI] and OH lines, revealing significant departures from LTE, indicating a much lower abundance of log(O/H)+12=8.66.
- Ar and Ne aren't seen in the Solar photosphere, so deduced from coronal material, relative to oxygen. The decrease in oxygen also causes Ar and Ne to be scaled down.



Consequences?

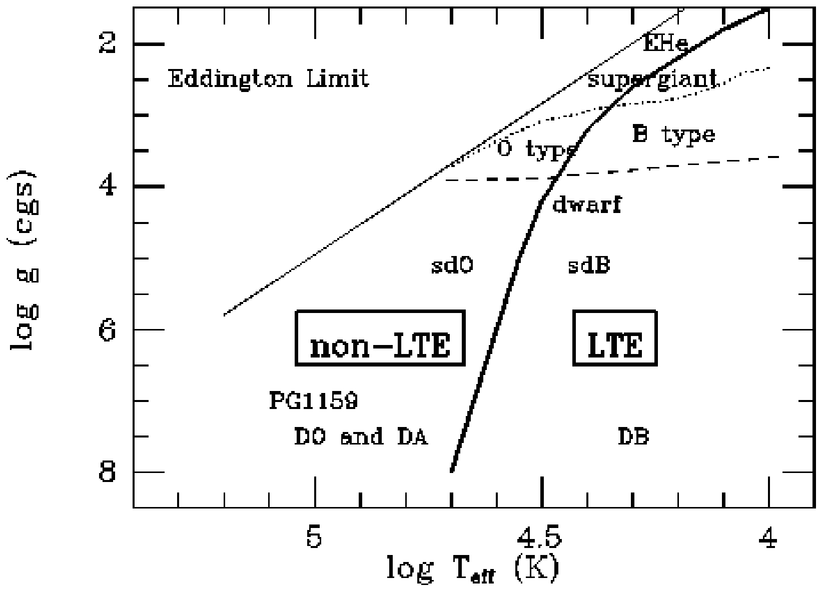
The Solar metal mass fraction **falls** from Z=0.019 to Z=0.013, reconciling some long-standing problems (e.g. agreement with local ISM abundances, e.g. Orion nebula), BUT there is now a helioseismology (sound speed, density below convective zone) discrepancy for the Sun, which can be reconciled in following ways:

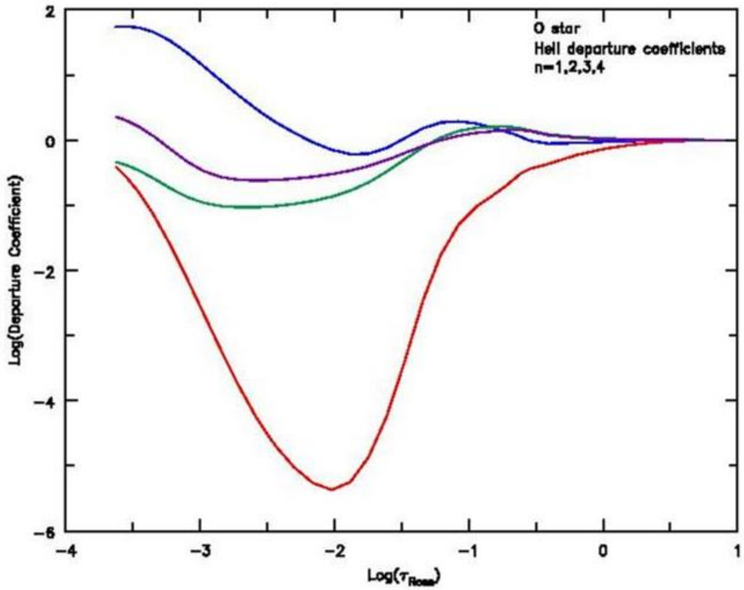
- Missing opacity from OPAL calculations? Need 7% at log T=6.4, though new OP calculations suggest <2.5% missing in OPAL.
- Problems with diffusion in interior models?
- Problems with abundance of Ne (indirectly inferred from Ne/O in solar corona). Needs factor 3 increase!

Overall, experience from Solar analysis suggests that determination of stellar abundances may be less certain than is normally considered!

Non-LTE for hot stars

Radiation field is so intense in hot stars (O-type, OBA supergiants, WDs) that their populations are only weakly dependent on local (T_{eff}, N_e), consequently LTE represents a poor assumption.

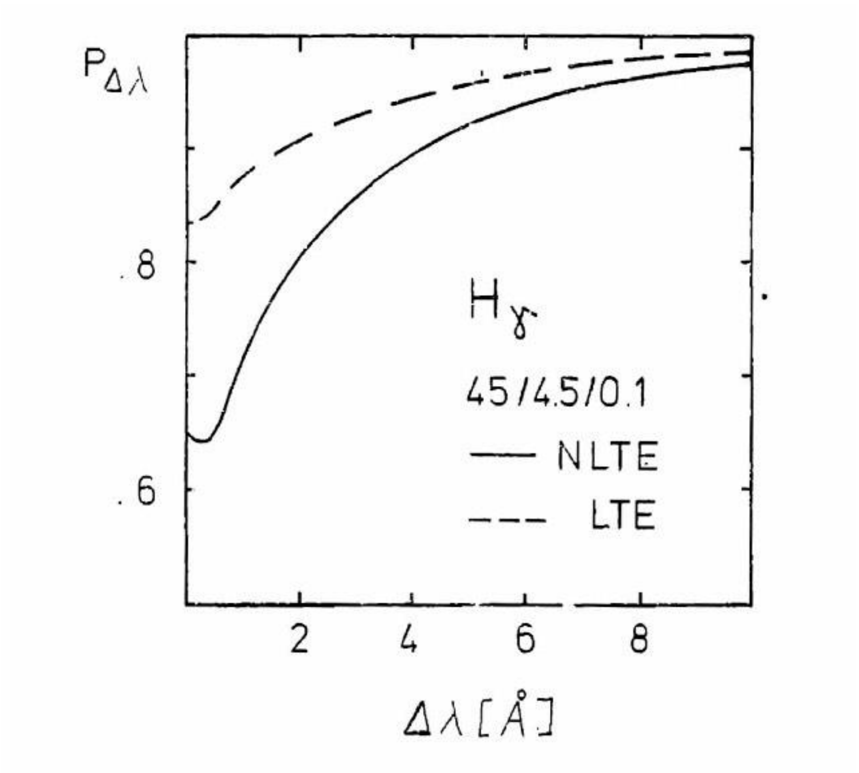


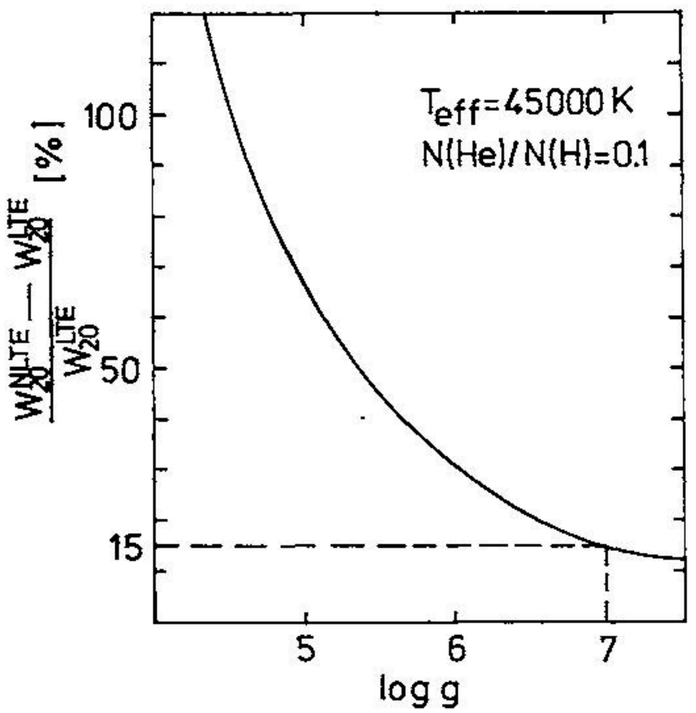


In O stars, LTE profiles are much too weak. Departure coefficients (non-LTE/LTE-pop) shown here for n=1,2,3,4 for HeII can differ greatly in wind and photosphere, making HeI & HeII lines *much* stronger.

LTE vs NLTE in hot stars

Difference between NLTE and LTE in Hγ line profile for an O-star model with T_{eff} = 45000K and log g = 4.5

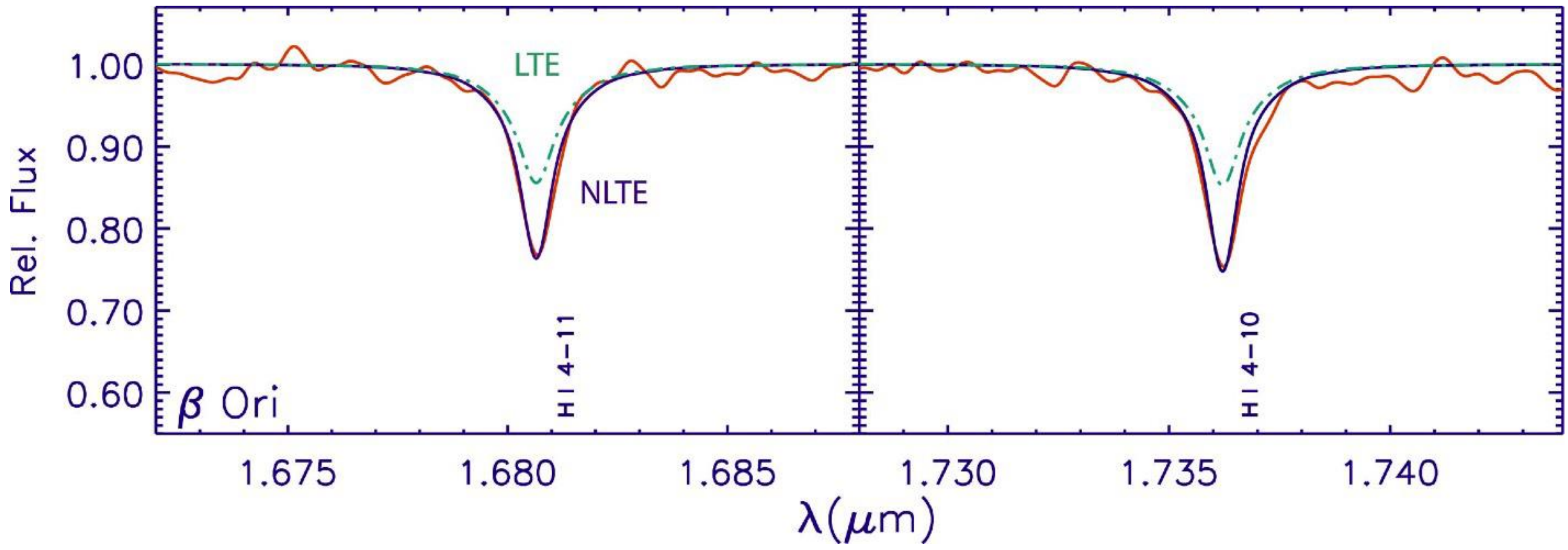




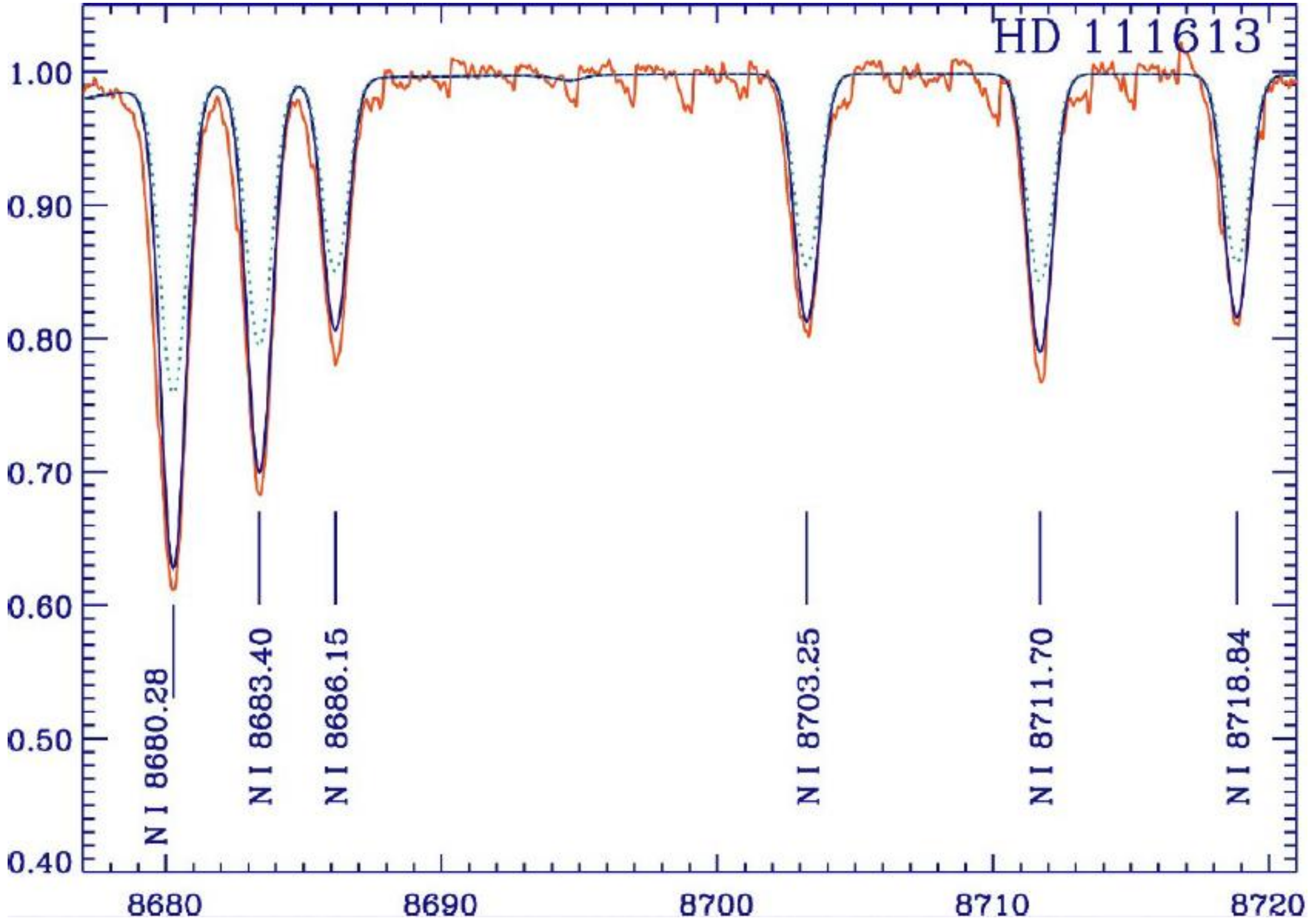
Difference between NLTE and LTE Hγ equivalent width as a function of log g for T_{eff} = 45,000 K for subluminous O stars

LTE vs NLTE: hydrogen lines in IR

Brackett lines



LTE vs NLTE: nitrogen lines



Non-LTE in OBA stars

- Hydrostatic equilibrium is invalid in OBA supergiants – their tenuous atmospheres lead to a drop in the line source function below LTE (Planckian) value.

- In the blue-violet spectra of B stars, some He I lines are formed in LTE, however red and IR lines are not collision dominated, instead photoionization-recombination processes dominate, so non-LTE is necessary.

- In A supergiants, reliable metal abundance determinations require non-LTE treatment – lines become stronger in non-LTE with corrections of up to factor of 10 for strong lines.

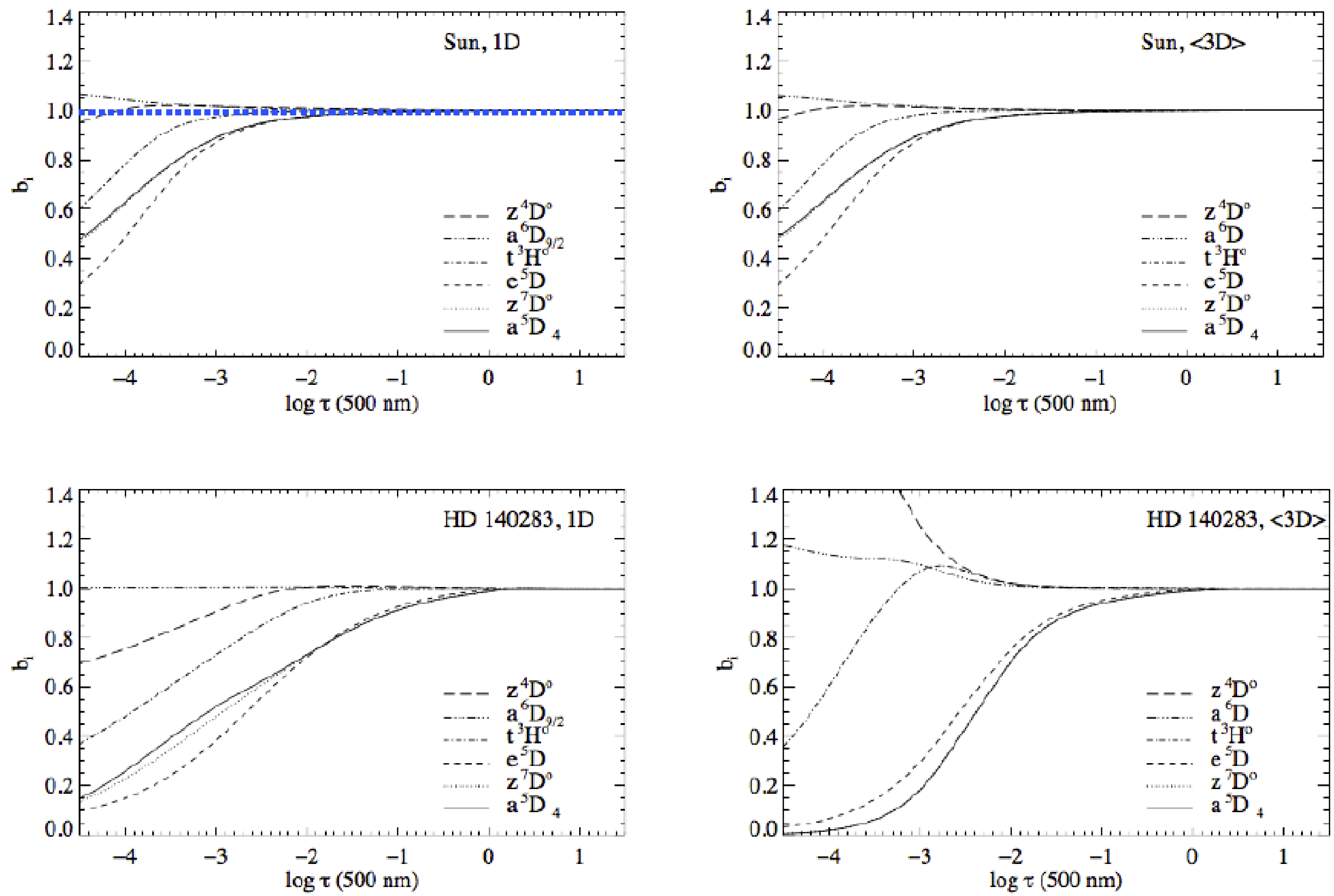
LTE vs NLTE in cool stars

LTE underpopulated

LTE overpopulated

LTE underpopulated

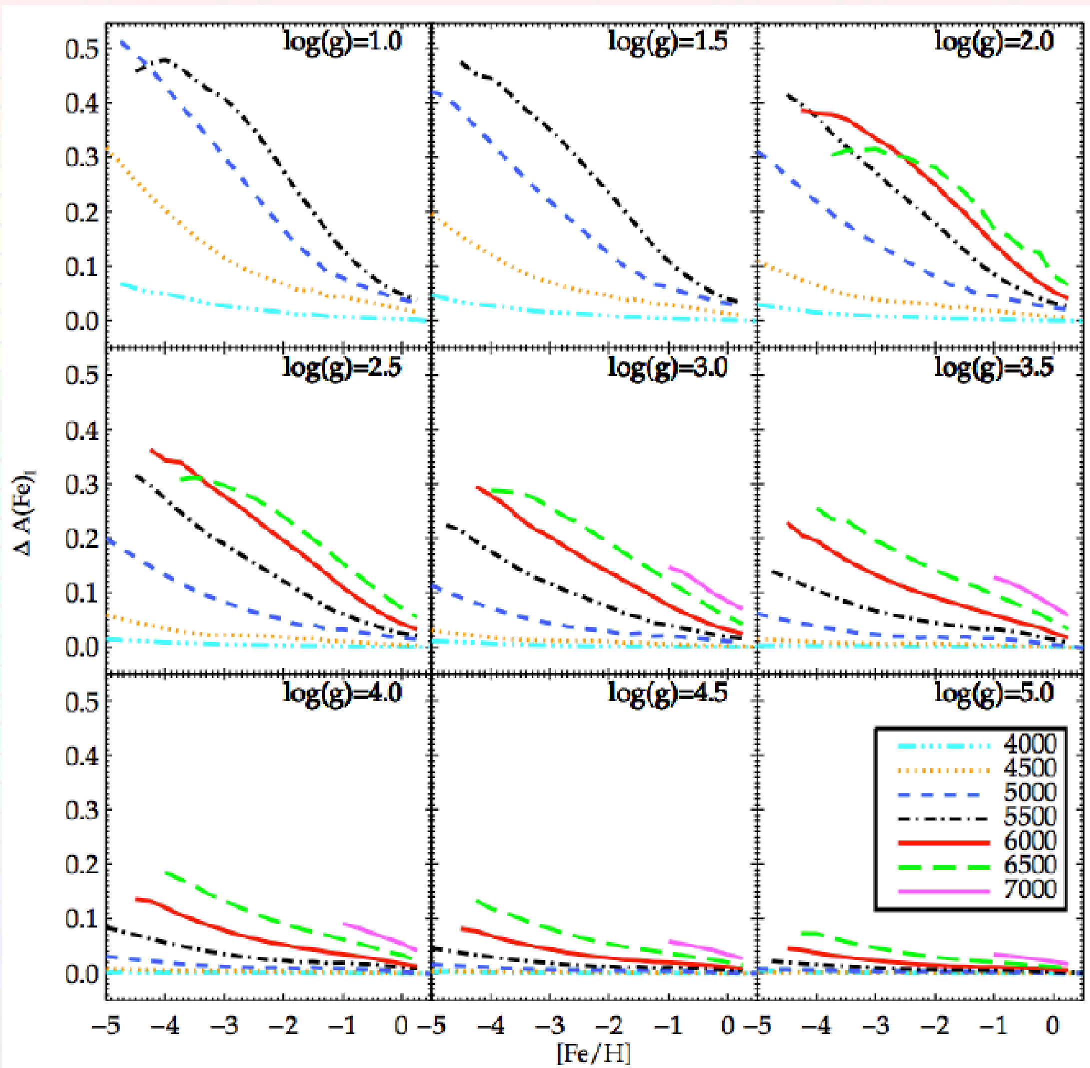
LTE overpopulated



NLTE effects & stellar parameters

decreasing [Fe/H] (less metal line blocking)

decreasing log(g) (collisional processes become less important)



increasing temperature (radiative processes become more important)

Lind et al. 2014

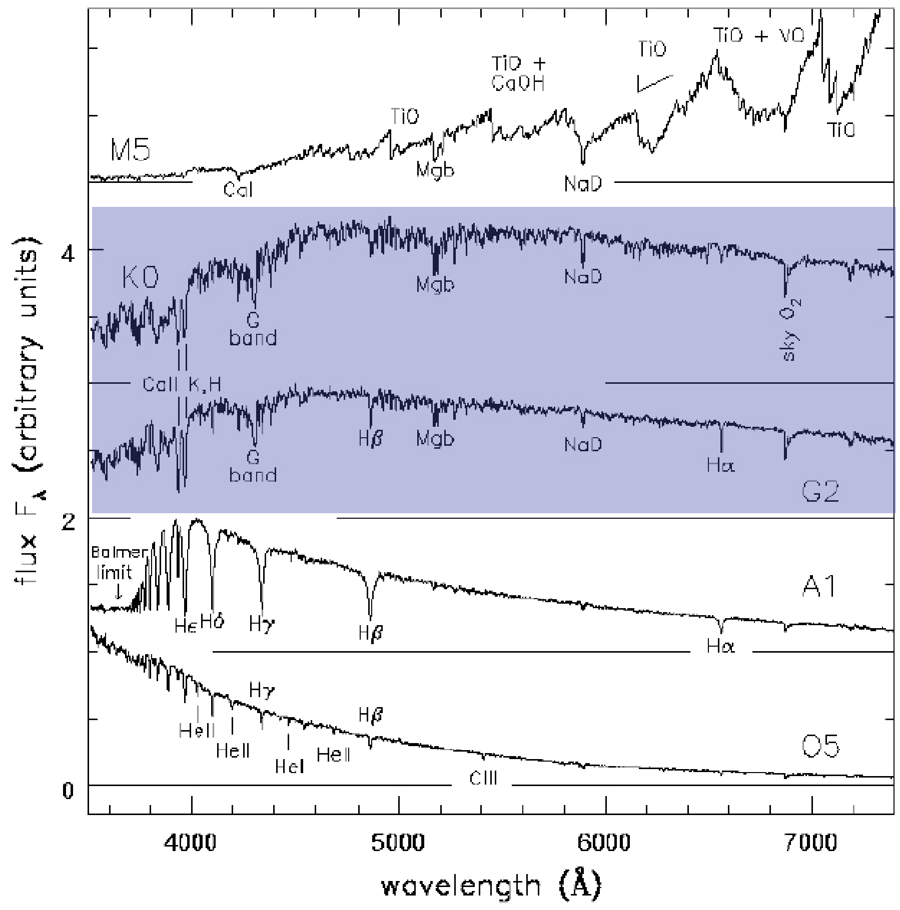
Summary

- If LTE does not hold, Saha-Boltzmann no longer describes excitation and ionization conditions – need to solve rate equations for statistical equilibrium – much more complicated!

- Non-LTE is necessary for hot stars, coronae of cool stars, M-type stars (as well as in nebulae and ISM).

Spectral type sequence

Spectral Types: temperature sequence



T ~ 4000K
Molecules!

Mainly neutral metal lines

T ~ 6000K

Ionised Metal lines

T < 11 000K

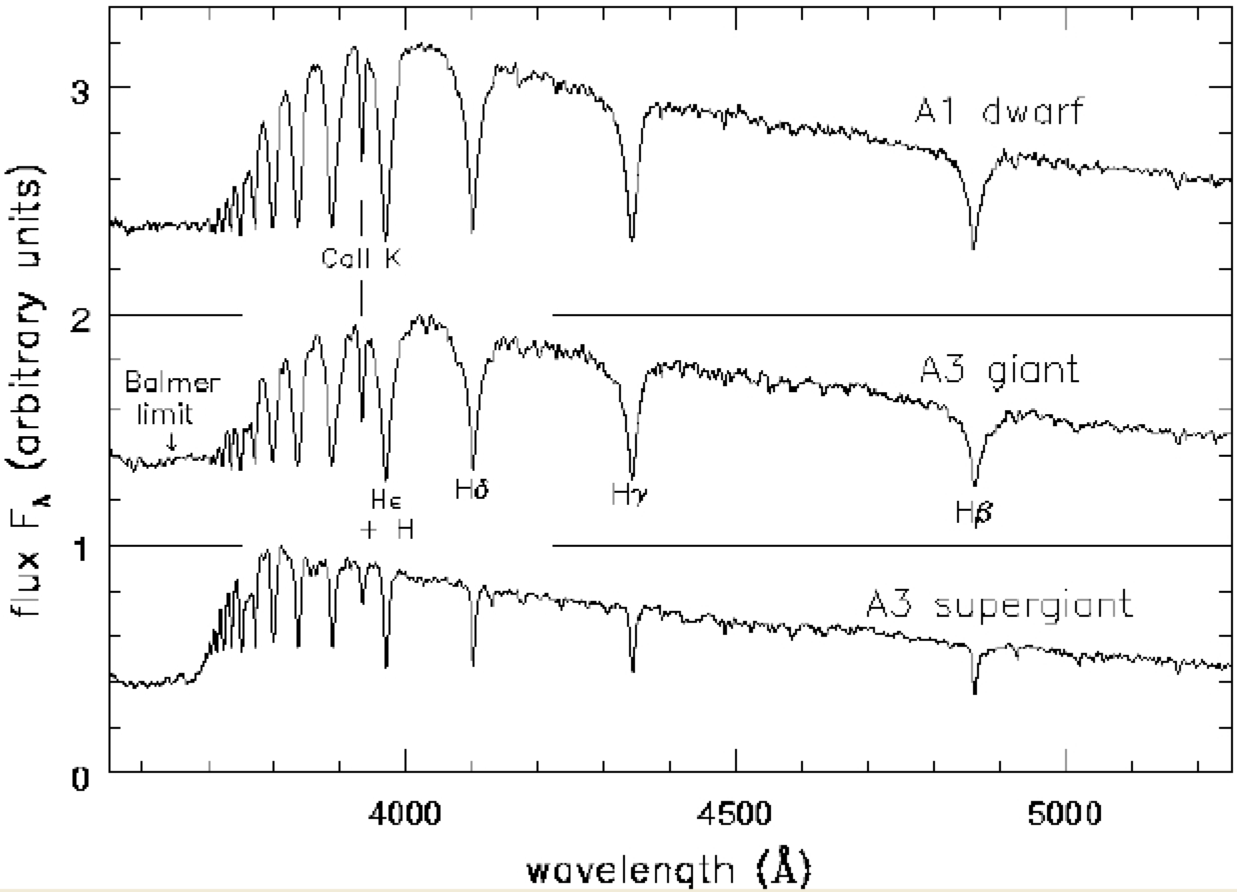
Dominated by neutral H

T ~ 30 000K

Highly ionised species

Line Broadenings

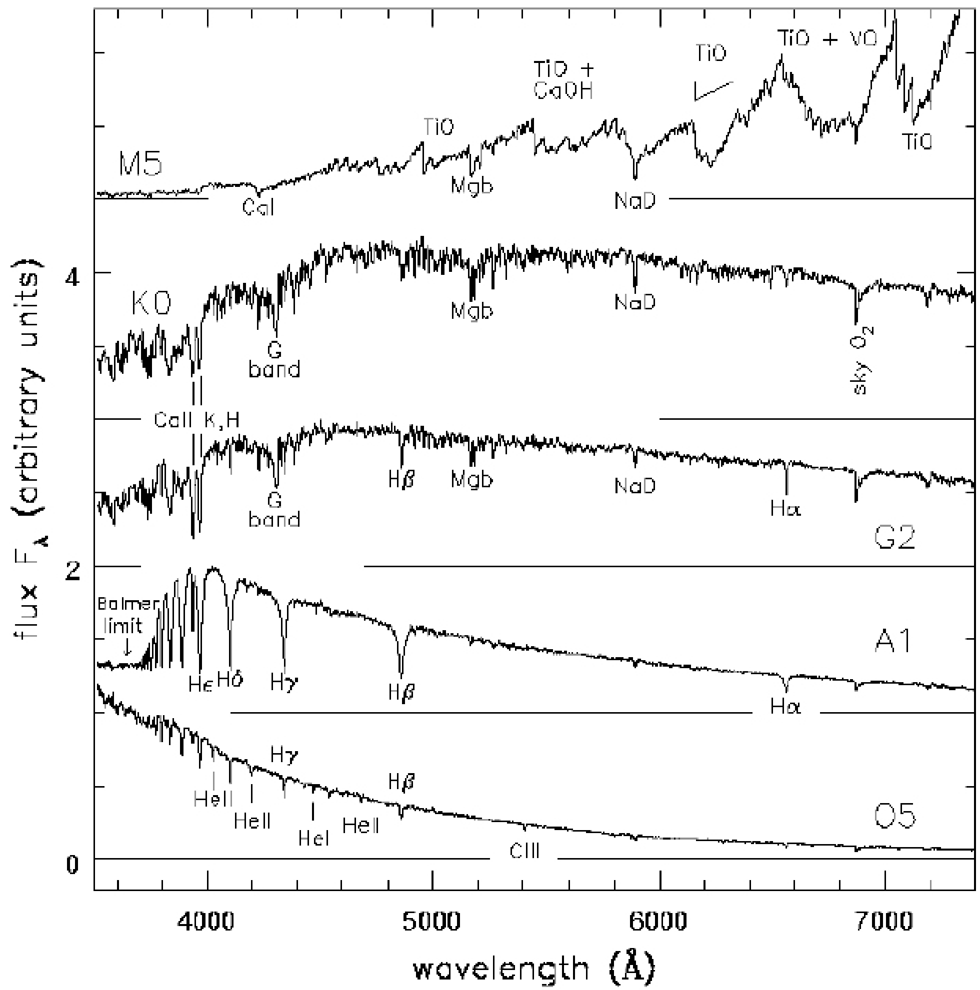
For example:
Stark Effect



He and Metals

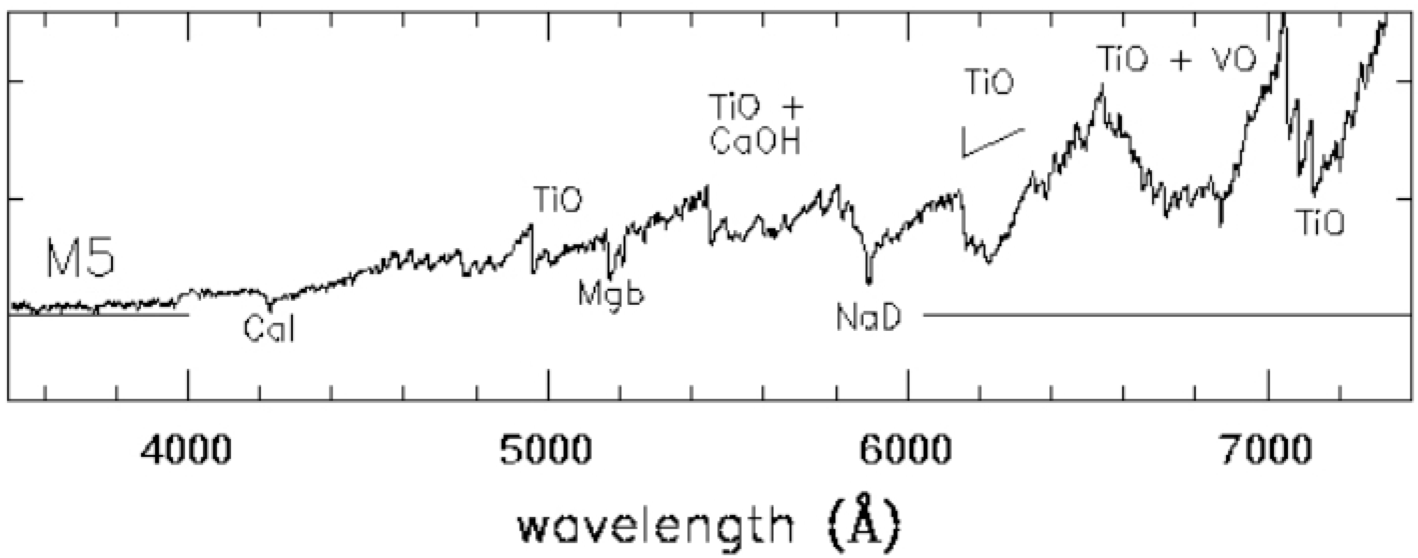
Metal are strongests when temperature is low enough that lower ionization stages are populated.
The metal lines become progressively stronger as the temperature cools and dominate in the F, G, K stars.

Helium is the second most abundant element, but only in the hottest stars (O and B) do He atoms show up in their excited levels where they can absorb visible light. For the very hottest O stars we also see HeII lines.

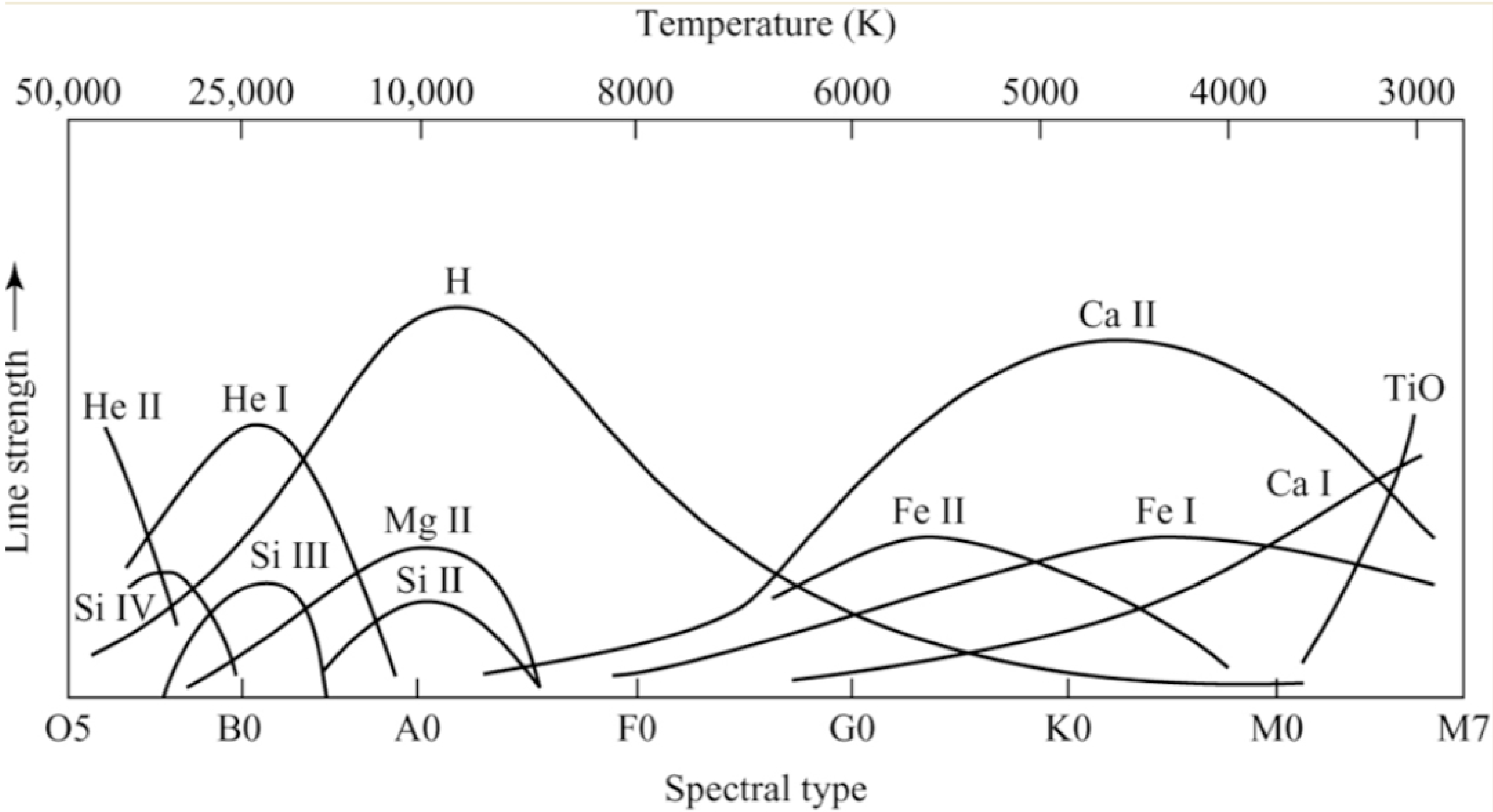


Molecular Bands

For very cool stars (M, L, T type) the atmospheres are sufficiently cool that simple molecules can form. These can absorb not only in electronic transitions, but also in vibrational and rotational modes. These create “bands” of absorption which can reduce the flux in vast portions of the spectrum. In M stars, TiO is a common important molecule. In L and T stars, other molecules such as CO, H_2O and CH_4 become important.



Relative Strength of Spectral Lines



Spectral classification