ASTRONOMICAL DATA ANALYSIS WITH COMPUTERS

Vitaly Neustroev

Practical sessions

- **Location:** YL124
- **Time:**

  Monday: 14:15 - 18:00
- **News, Lectures and Problems:** http://vitaly.neustroev.net/teach/spring-2017.html

Text Books

- **My course "Observational Astrophysics and Data Analysis"** *http://vitaly.neustroev.net/teach/autumn-2016/*
- *Astrophysical techniques* (5th Edition - 2008) – C.R. Kitchen: Taylor & Francis / CRC Press. ISBN 978-1-4200-8243-2.
- *Observational Astrophysics* (3rd Edition – 2012) – P. Léna, D. Rouan, F. Lebrun, F. Mignard, D. Pelat, Translated by S. Lyle: Springer. ISBN 978-3-642-21815-6.
- An Introduction to Astronomical Photometry Using CCDs (Oct. 22, 2006) – W. Romanishin: the latest version of this book can be downloaded from http://vitaly.neustroev.net/teaching/2016b/wrccd22oct06.pdf
- http://vitaly.neustroev.net/teaching/2017a/iraf.pdf
- http://vitaly.neustroev.net/teaching/2017a/irafmanual.pdf

Methods of Observations

- Imaging
- Photometry
- Spectroscopy

The Primary Tools of Astronomy

Light is the only thing we can work with. What can we do with it?

- Imaging – we can take "pictures" of the things we see. But pictures alone tend to lack the "quantitative" aspect that is needed for most serious scientific studies.
- Photometry – the technique that measures the relative *amounts* of light in different wavelength ranges.
  But these ranges are too wide to provide detailed information on the light's spectral distribution.
- Spectroscopy (spectrophotometry) – the most informative technique of light analysis, that measures how much light an object produces at various wavelengths of light.

Imaging

- Mapping the distribution of celestial sources on the sky in order to locate the position of source precisely – astrometry.
- Getting information on the source's form and that of its local environment.

Primary reduction and Data Analysis

- **Reduction** is the process of turning raw data into a calibrated product; anything after that is **Data Analysis.**
- Data you obtain from observations must be reduced; data you download from the archives may be partially reduced and calibrated.

Charge Coupled Device detectors

- CCDs – a replacement for photoplates. This was an astronomer's dream for decades.
- Ubiquitous in optical and very near infra-red astronomy
- The detector of choice at wavelengths from 4000 Å to 10000 Å.
- Robust enough to fly on space missions.

How does a CCD work?

- In order to produce an image, a CCD must accomplish four functions:

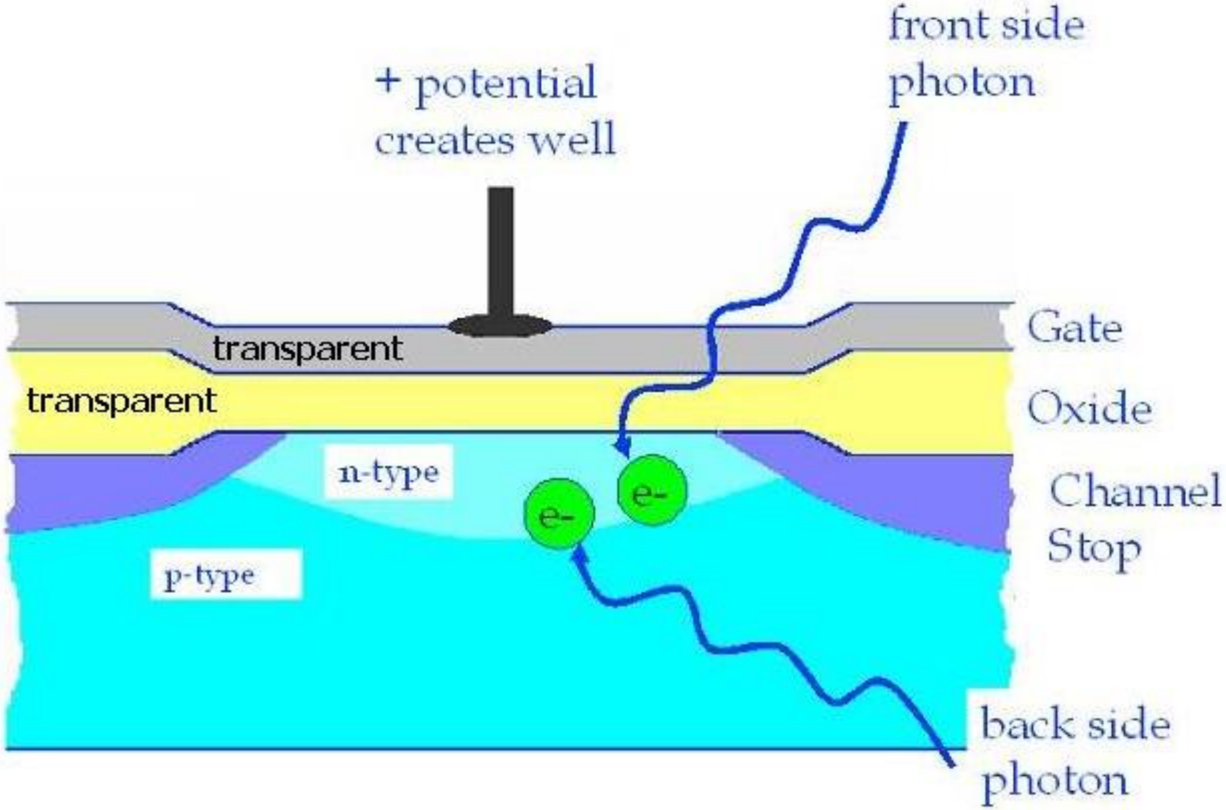
1) generate photoelectrons
2) collect electrons
3) transfer the collected charges
4) read the charges

How does a CCD work?

- The first function is based on the photoelectric effect.
- A CCD is a silicon wafer which is exposed to radiation.
- The light absorption in the silicate network of the CCD generates these photoelectrons, in proportion to the number of incident photons.
- The latter are immediately collected in “picture elements” so-called pixels, closest to where the photons fell on the chip. The picture elements (pixels) of the CCD are defined by the electrode structure which is applied to this wafer.
- The electrodes form potential wells to prevent the collected charges from escaping.

How does a CCD work?

Photoelectrons are collected deep in the well beneath the positive potential set up by the gate electrode, isolated from the surface by the n-type silicon buried channel.

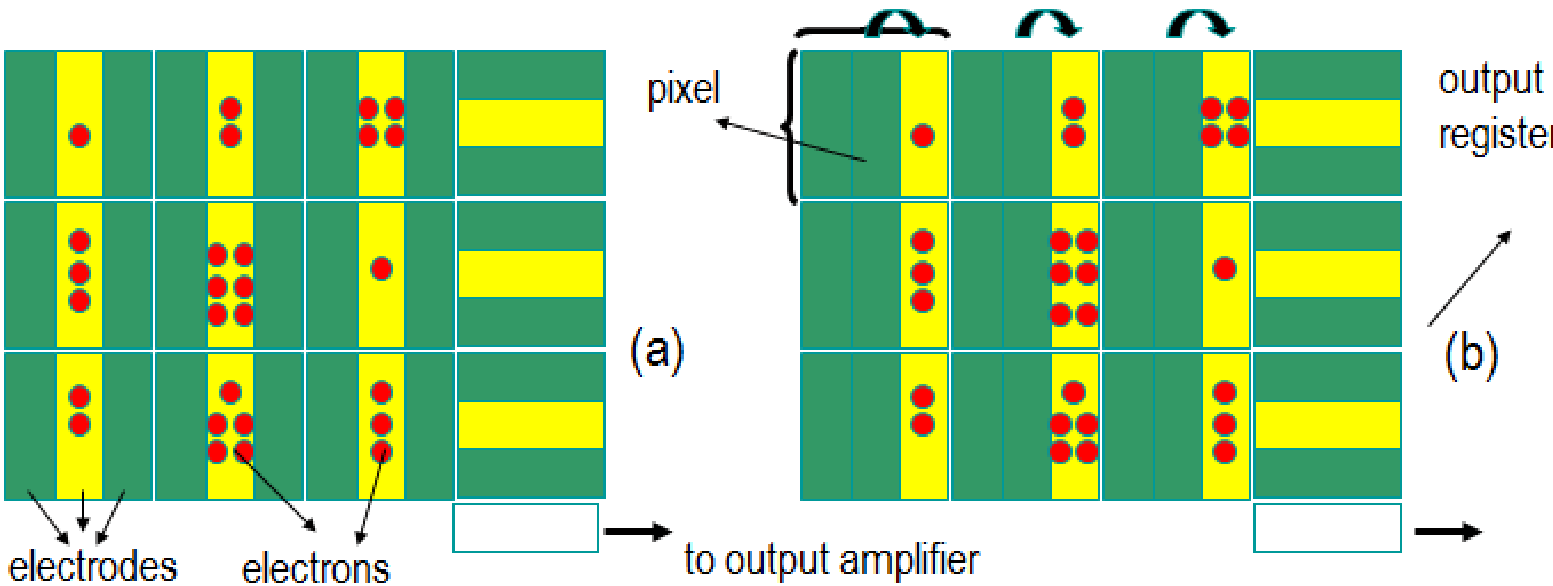


How does a CCD work?

- As you expose the CCD to radiation, electron-hole pairs are generated and the electrons build up in the electron storage areas immediately below the positive potential electrodes.
- After a while (seconds to minutes) the CCD contains an electrostatic representation of the pattern of incident radiation on it.
- This somehow must be read out and stored in digital form.

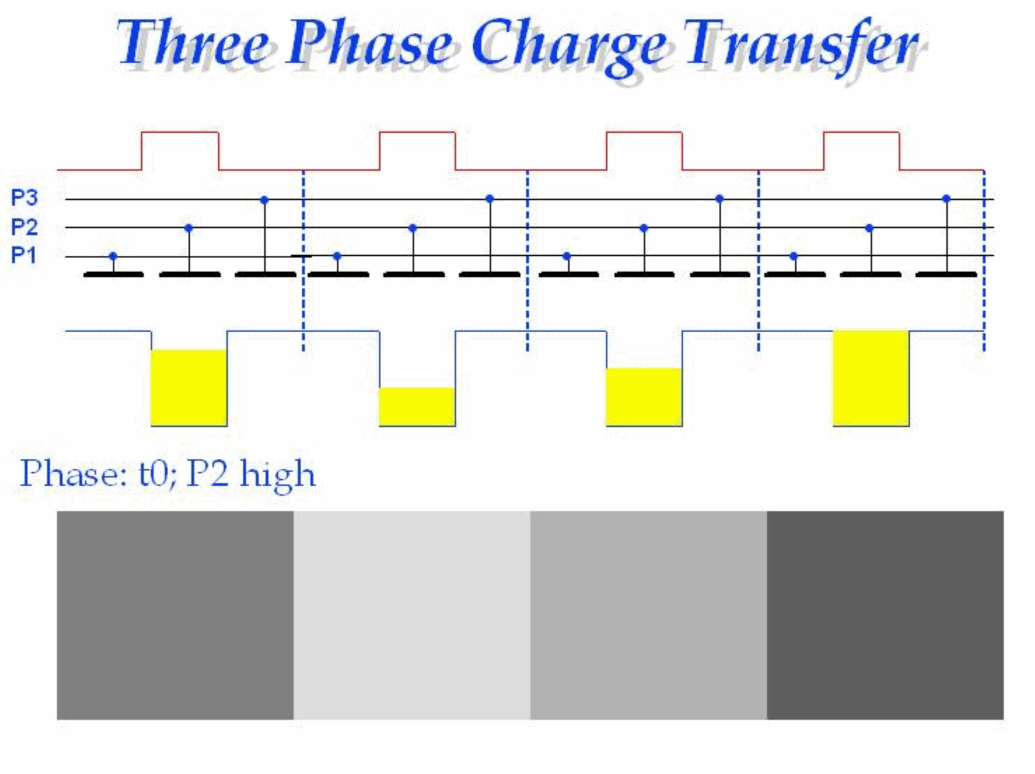
How does a CCD work?

- Each pixel is divided into 3 regions (electrodes who create a potential well). For the charge collection process during an exposure the central electrode of each pixel is maintained at a higher potential (yellow) than the others (green).



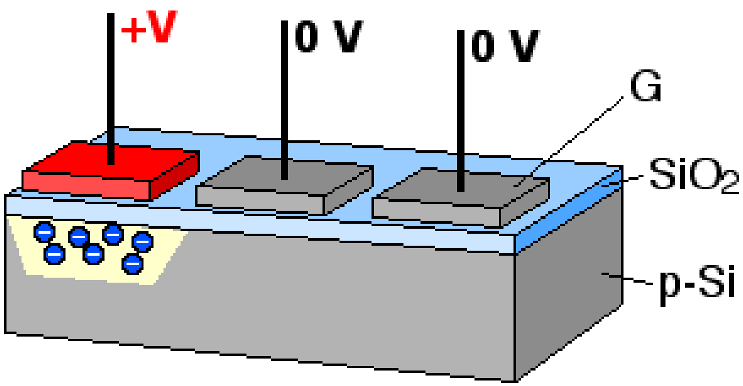
How does a CCD work?

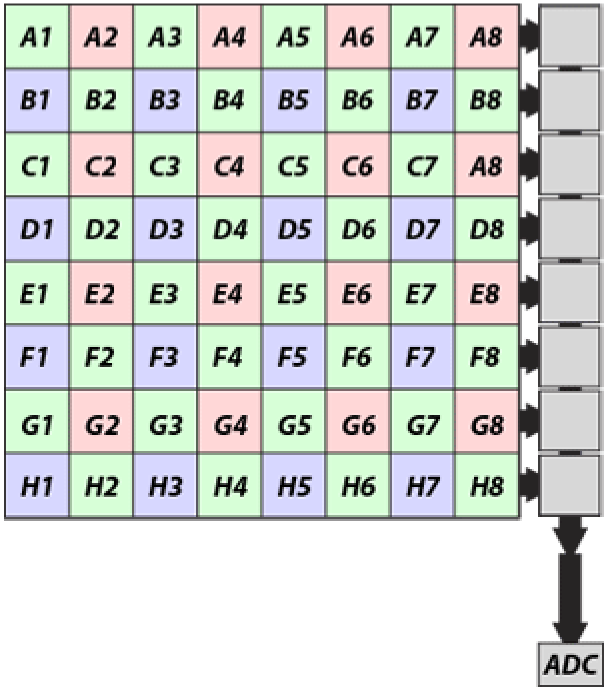
- By changing the potential of the electrodes in a synchronized way, electrons are transferred from pixel to pixel. Charges on the right are guided to the output register



How does a CCD work?

- The readout register is shifted to the right by one pixel, and the pixel at the bottom right is shifted into a readout capacitor.





CCD readout

- This process is then repeated until each pixel in the readout register has been digitized.
- The image is then shifted right horizontally by one more pixel, the next column is shifted into the readout register, and this is digitized in the same way.
- The whole process is repeated until the entire image is read out.
- For a 2048 x 2048 pixel CCD it takes approximately 10-60 seconds to read out the whole chip.
- A CCD read out this way is a line transfer CCD.

CCD readout

- The channel stops between rows are permanent as charge does not move vertically except in the readout register.
- These channel stops are biased to negative potential by doping, hence charge cannot leak horizontally.

Advantages of CCDs

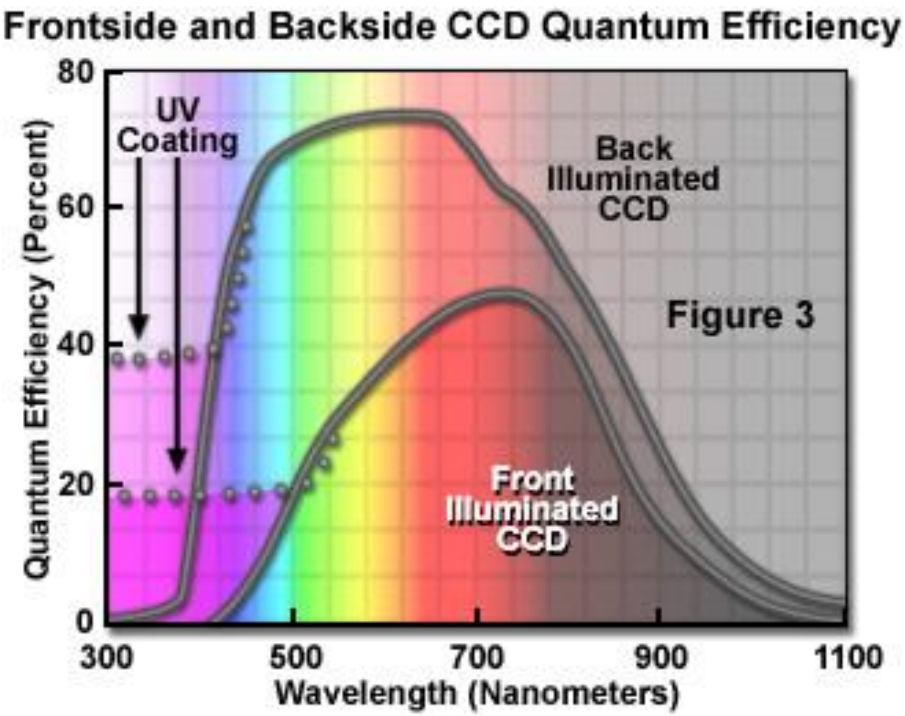
- Good spatial resolution
- Very high QE of up to 80%.
- Large spectral window
- Very low noise
- High photometric precision
- High dynamic range
- Very good linearity
- A reliable rigidity (no physical distortion, etc)

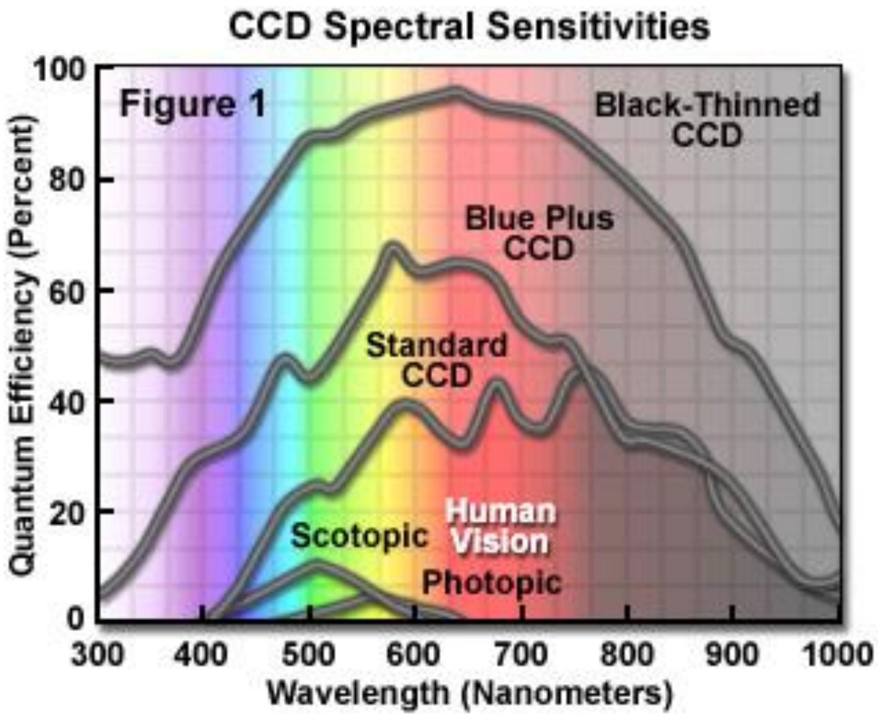
Advantages of CCDs

- **Good spatial resolution:**
- Today, most common CCDs have 2048 × 2048 pixels. But there exist even larger CCDs with 4096 × 4096 pixels or 4096 × 8192 pixels (10 k x 10 k will be build soon).
- For realizing even larger chips (and since larger CCD chips are very expensive), several small chips can be placed together resembling a CCD mosaic.

Advantages of CCDs

- Quantum Efficiency:





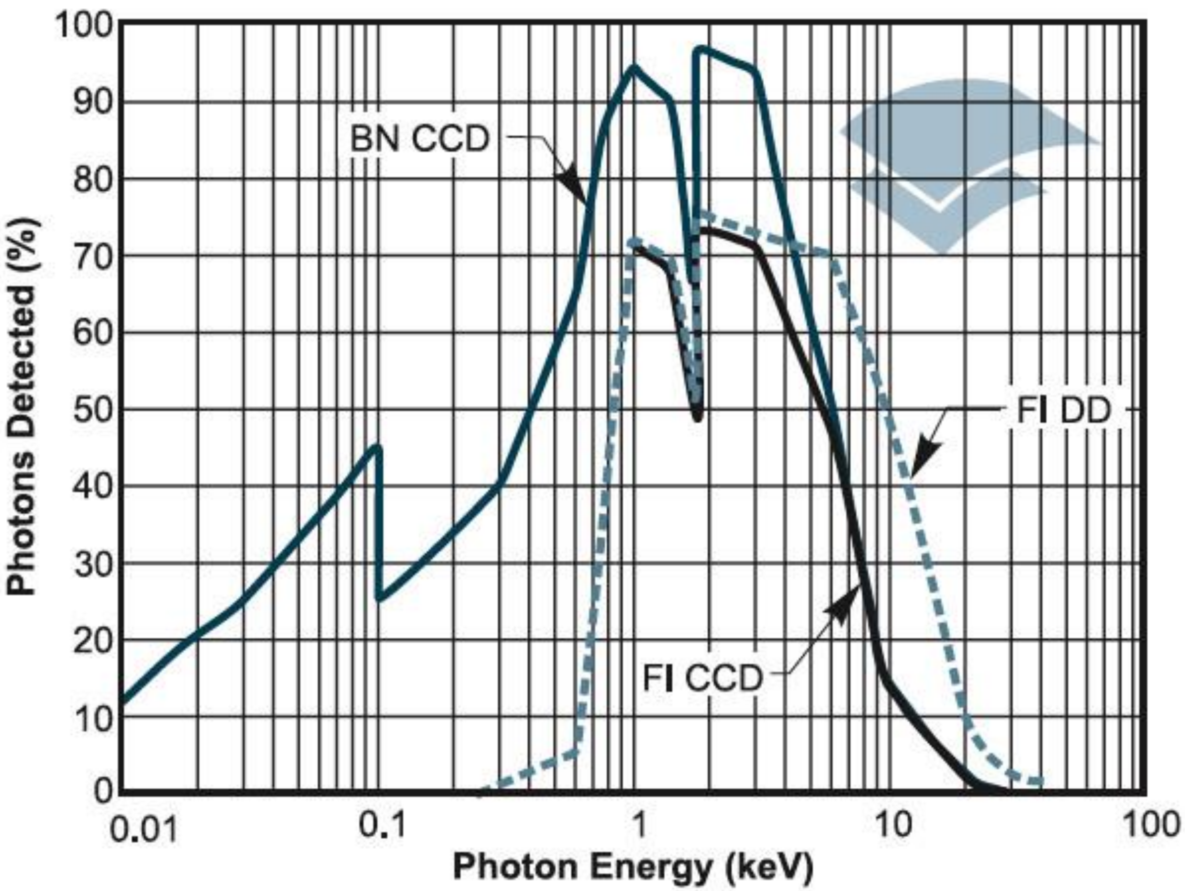
Advantages of CCDs

Spectral Range

FI: front illuminated

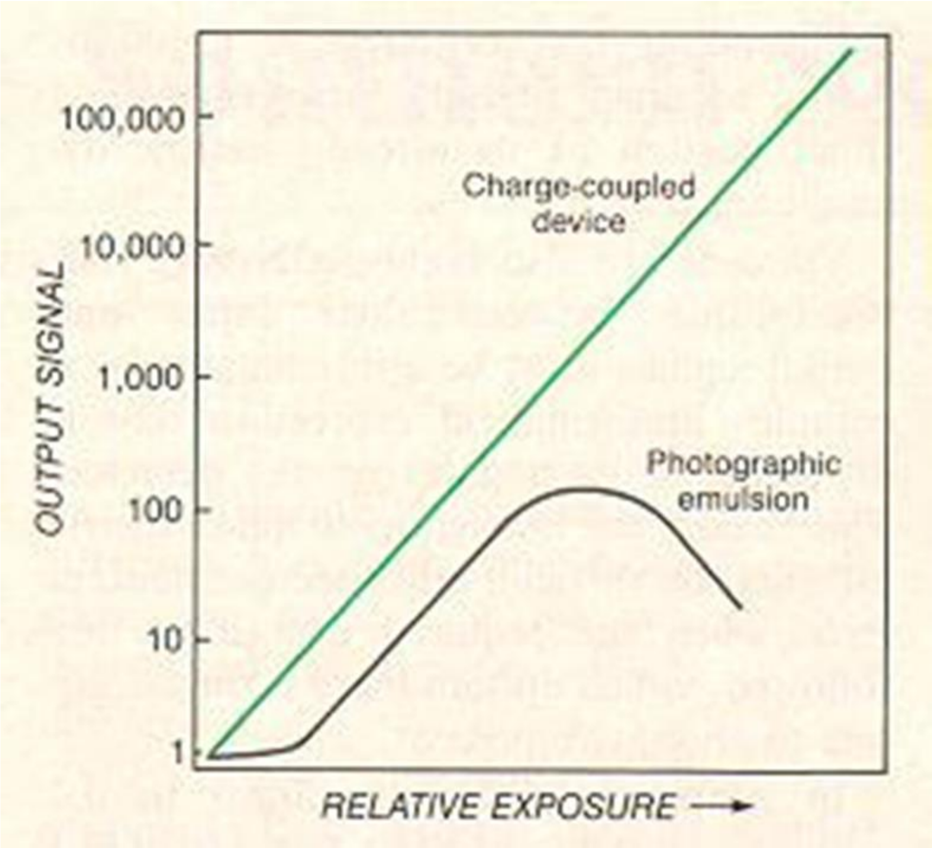
BN: back illuminated, no coating

DD: deep depletion



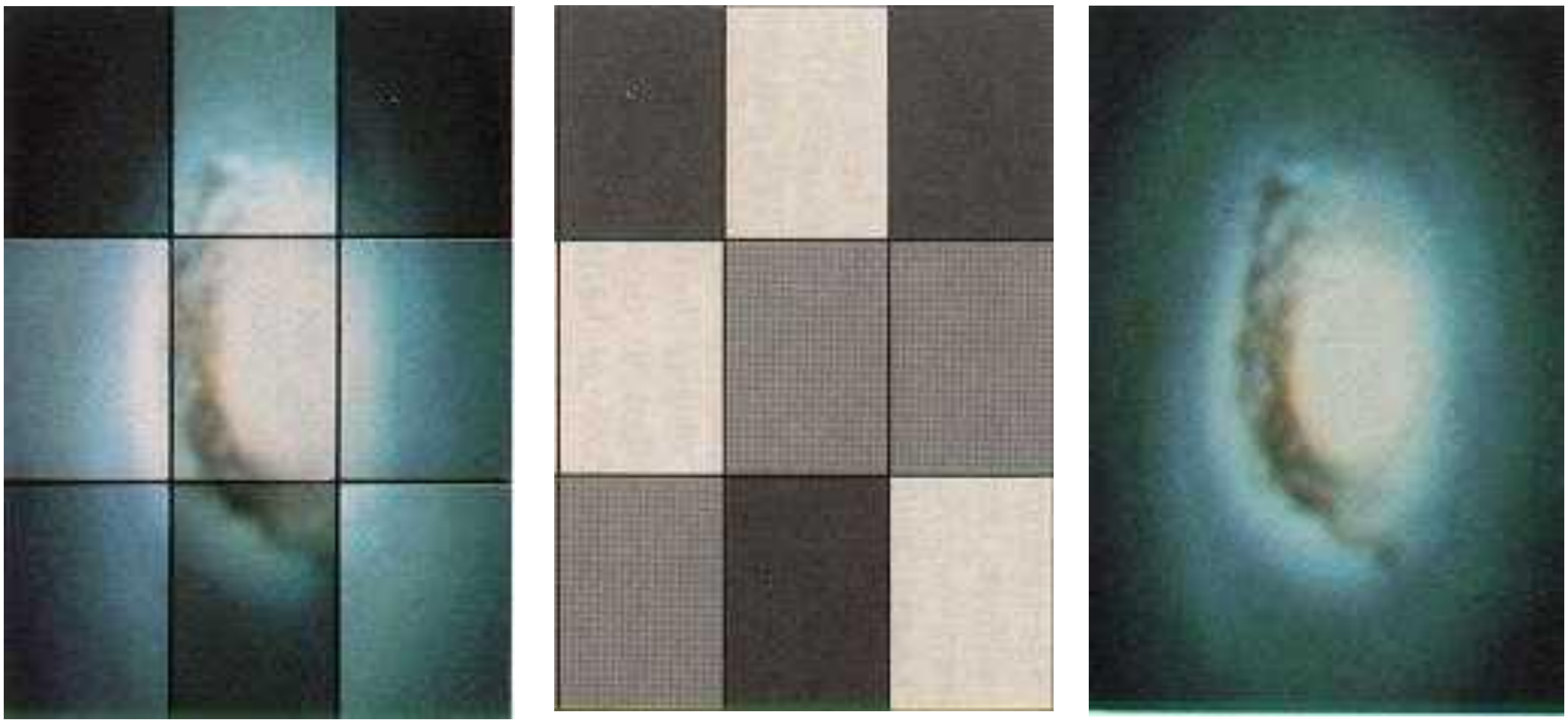
Advantages of CCDs

- **Linearity and Dynamic Range:** CCDs are extremely linear detectors. Therefore CCDs enable the simultaneous detection of both very faint and very bright objects. The dynamic range of CCDs is about 100 times larger compared to photoemulsions.



Advantages of CCDs

- Flat field technique.



Disadvantages of CCDs

- **Size:** The size of a single pixel is in the order of 8 × 8, 15 × 15 or 25 × 25 microns. Therefore the size of CCD chips remain quite small, especially by comparing CCDs to classical photographic plate images. E.g., a CCD with 2048 × 2048 pixels of 15 microns measures only 3 × 3 cm^2. In contrast a photographic plate for a Schmidt telescope can be as big as 30 × 30 cm^2, equivalently to a CCD chip with 400 million pixels!

Disadvantages of CCDs

- **The dark current** is background signal generated by thermal effects. Because of the dark current CCDs are run cooled, to reduce the possibility of thermal excitation of electrons across the band gap.
- CCDs are operated at temperatures of around 140K, to reduce thermal effects.
- Dark current at 140K is typically 10^{-4} electrons/s/pixel, i.e. negligible.

Disadvantages of CCDs

- **Cosmic rays, X-rays, and particle radiation:**
- There are a number of types of radiation which can interact with the silicon to produce several tens of electron-hole pairs in a cluster, which appears as a bright spot (if the radiation is normal to the detector) or a streak if it is steeply inclined. These radiation events are:
  - Secondary muons in cosmic ray air showers.
  - X rays emitted by UV transmitting glass in the optics of the instrument.
  - Radioactivity from heavy metal impurities in the cryostats.
- These events are identified, classified and rejected by splitting the CCD exposure into two or more equal parts, the hits don't occur in the same place.

Disadvantages of CCDs

- **Saturation** (not a real problem): Typically the full well capacity of a CCD pixel 25 µm square is 500,000 electrons. If the charge in the well exceeds about 80% of this value the response will be non-linear. If it exceeds this value charge will spread through the barrier phase to surrounding pixels.
- This charge bleeding occurs mainly horizontally, as there is little vertical bleeding because of the permanent doped channel stops.
- Readout register pixels are larger, so there is less saturation effect in the readout register.

Disadvantages of CCDs

- **Charge Transfer Efficiency:** When the wells are nearly empty, charge can be trapped by impurities in the silicon. So faint images can have tails in the horizontal direction.
- Modern CCDs can have a charge transfer efficiency per transfer of 0.9999995, so after 2000 transfers only 0.1% of the charge is lost.

CCDs: readout noise

- CCDs suffer from **readout noise** which has a variety of sources:
  - The output Field Effect Transistor. This is the ultimate limit to the readout noise, at a level of 2-3 electrons.
  - Transfer loss fluctuations. During transfer an amount of charge is left behind, but this amount varies. Transfer noise is given by: $\sigma_{tr} = \sqrt{(2\zeta n N_0)}$ where ζ = 1-CTE is the fraction of charge not transferred, n is the number of transfers and N_0 is the original charge. For faint sources (≈100 electrons) this noise is less than 1 electron.

CCDs: Other noise sources

- **Fixed pattern noise.** The sensitivity of pixels is not the same, for reasons such as differences in thickness, area of electrodes, doping. However these differences do not change, and can be calibrated out by dividing by a flat field, which is an exposure of a uniform light source.
- **Bias noise.** The bias voltage applied to the substrate causes an offset in the signal, which can vary from pixel to pixel. This can be removed by subtracting the average of a number of bias frames, which are readouts of zero exposure frames. Modern CCDs rarely display any fixed pattern bias noise

CCDs: Interference Fringes

- In thinned CCDs there are interference effects caused by multiple reflections within the silicon layer, or within the resin which holds the CCD to a glass plate to flatten it.
- These effects are classical thin film interference (Newton's rings).
- Only visible if there is strong line radiation in the passband, either in the object or in the sky background.
- Visible in the sky at wavelengths > 700nm.
- Corrected by subtracting off a scaled exposure of blank sky.
- Fringing can dominate the noise in the redder photometric bands, or in narrow bands, and can sometimes force us back to using thick CCDs despite the loss in QE.

Image processing

- A CCD image is built up from three signal sources:
  - object: photons imaged onto each pixel of the CCD
  - dark current: thermal electrons collected in each pixel during the exposure
  - bias: a low level electrical signal added to each pixel during image readout

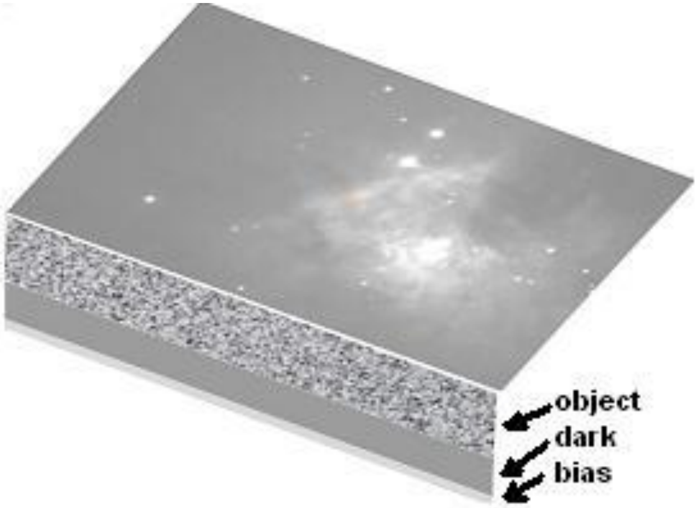
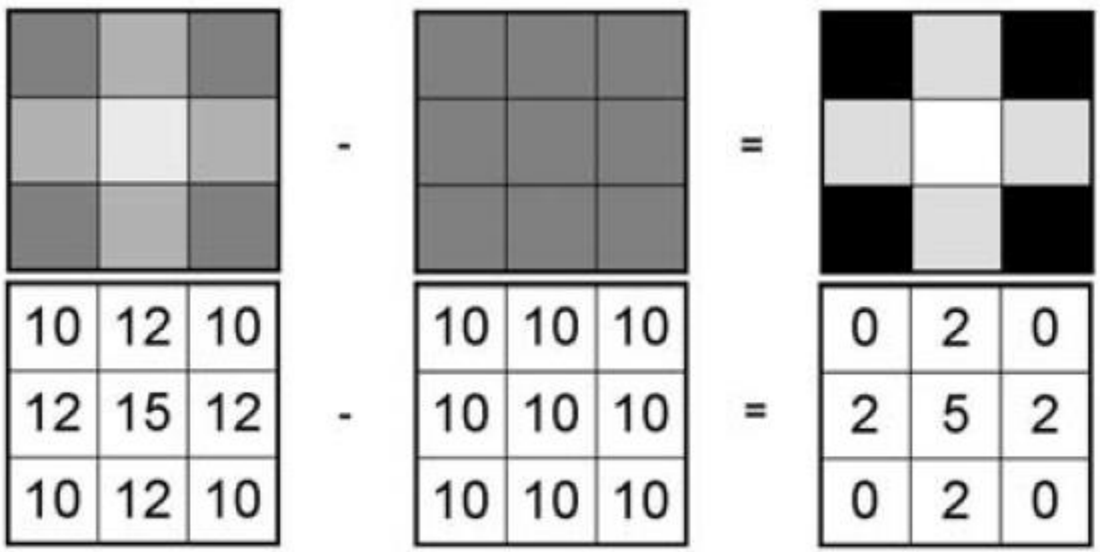


Image processing

- Most processing steps must be carried out pixel by pixel



A simple operation on digital image data to subtract one frame from another.

Image processing

- **Darks:**
  - In order to remove the accumulated background due to thermal dark current, it is usual to take several "darks" which are images taken with the CCD camera shutter closed. These darks should be of the same exposure time and camera temperature as the object exposure to be dark subtracted.
  - It is best to take a set of dark images and then to combine them to get a "master" dark based on the average of the dark set.
    [Note: it is usual to "median combine"]

Image processing

- **Bias:**
  - In addition to thermal noise, each pixel charge will carry with it a fixed offset voltage value called the bias. Thus even if the output coming from the CCD were exactly zero electrons for every pixel, there would still be a signal that would vary from pixel to pixel in a repeatable fashion. An bias frame is one taken to determine this bias pattern.
  - Again, it is best to take a set of bias images and then to combine them to get a "master" bias based on the average of the bias set.

Image processing

- **Flat Fielding:**
  each pixel on the CCD may have a different sensitivity to incoming photons due to small variations in individual pixel dimensions and quantum efficiency. For precision photometry it is necessary to calibrate such pixel-to-pixel variations and this is the function of "flat fielding". There are several different approaches but the two that are applicable to most observatories are
  - Dome Flats: A uniformly illuminated target is installed in the dome or attached to the front of the telescope
  - Sky Flats: Images are taken of the sky which is assumed to be uniform in brightness over the (usually small) field of view of the CCD

Image processing

- **Flat Fielding:**
  - Take an exposure of a source which will uniformly illuminate each pixel of the CCD
  - Pixel sensitivity (and perhaps other effects) will be a function of wavelength so separate flat fields are needed for each filter
  - Use a long enough exposure time to fill pixels to more than 50% of their full well capacity

Image processing

- **Flat Field** types:
  - Twilight Flat: Exposures of the twilight sky (well away from horizon). Can correct for all of the types of sensitivity variations but the twilight sky is typically much bluer than the typical program object.
  - Dome flat: Exposure of a special target, usually mounted on the dome (such a target is completely out of focus and thus effectively uniform). Does well on Pixel to Pixel variations, poorly on vignetting.
  - Sky Flat: Median combination (to remove stars) of many exposures of the night sky
    Good correction of vignetting, poor correction of Pixel to Pixel variation. Hard to get enough photons on CCD – exposures may be very long.

Image processing

- **Flats:**
  - Take a number of dome or long exposure (5 to 10 seconds) twilight flats
  - Do Dark and Bias corrections for each image
  - Average or median combine to get high S/N image
  - Make a copy and smooth with a large area filter (25 x 25 pixels) to remove pixel to pixel variations
  - Divide high S/N image by smoothed image
  - Normalize to 1.00 at the center
  - Save as a 32- bit real image
- The result is your "pixel flat"

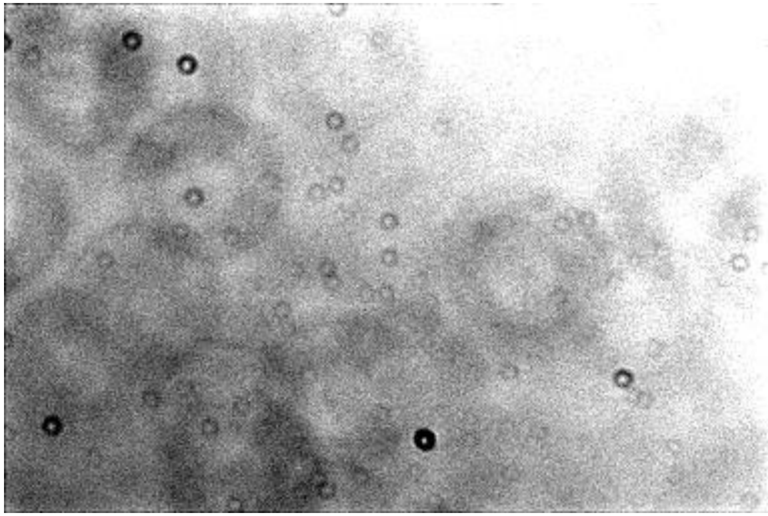


Image processing

- Cleaned image after dark, bias, and flat-field corrections

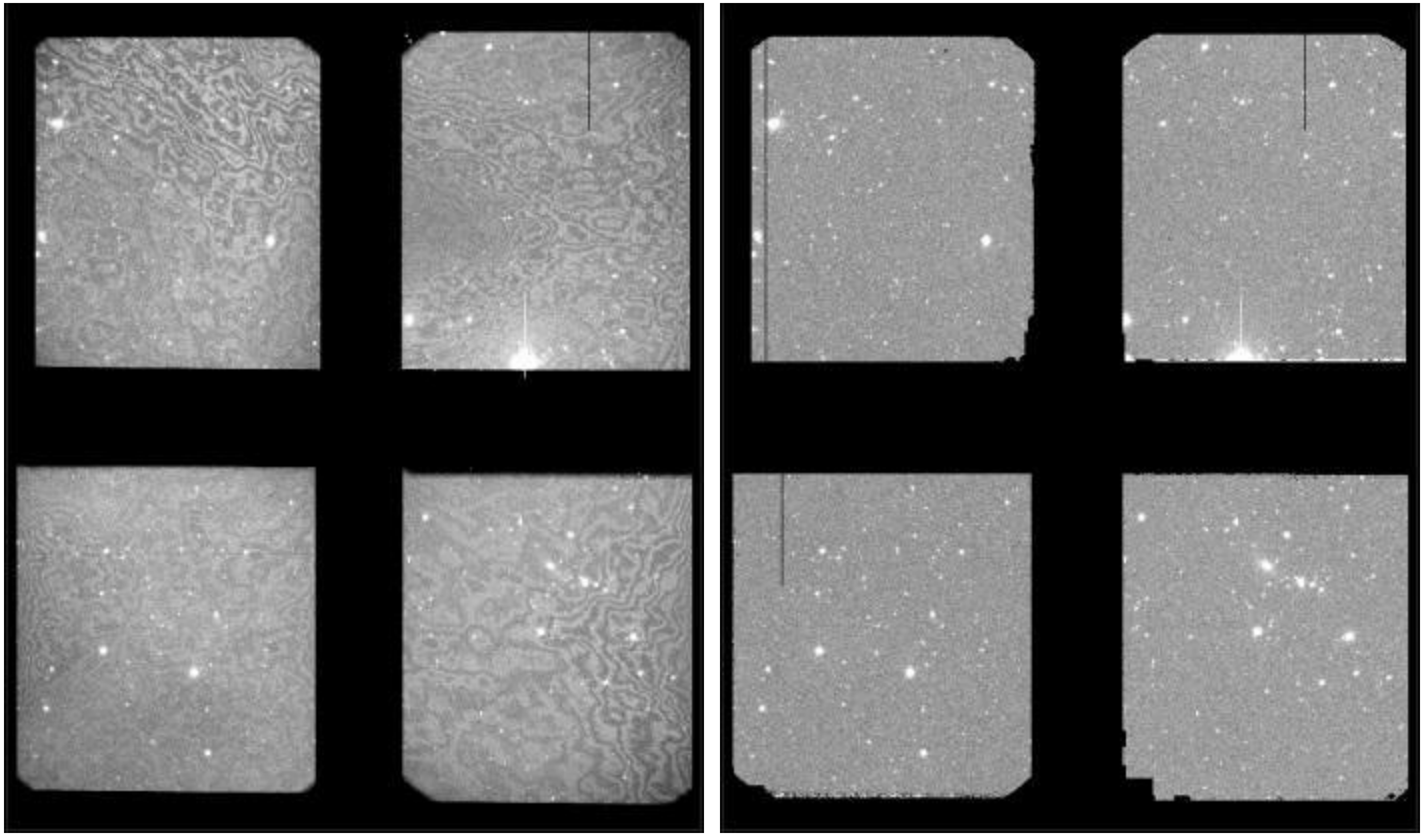


Image processing

- **Summary:**
  - Start with raw image
  - Subtract dark and bias images
    - use high s/n "master bias" and "master dark" frames for best results
  - Divide by flat field
    - use high s/n "master flats" for best results
    - flats must have dark and bias removed

Image processing

- **Do not forget: Cosmic Ray Events:**

Image processing

- There are several major computer packages:
  - IRAF
  - MIDAS
  - IDL (commercial)
  - Starlink